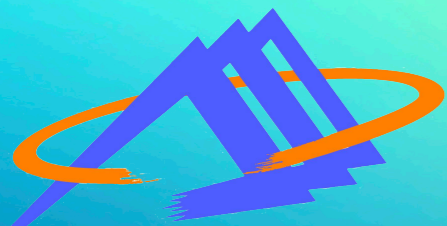


# Charge exchange X-ray emission in Spiral galaxies



Purple Mountain Observatory  
CAS, Nanjing, China

Shuinai Zhang (snzhang@pmo.ac.cn)

Daniel Wang (UMass) Randall Smith, Adam Foster (CfA)

Kinwah Wu (Mullard) Yiheng Chi (NJU)

Li Ji, Hang Yang, Huiyang Mao (PMO)

Charge exchange X-ray Universe

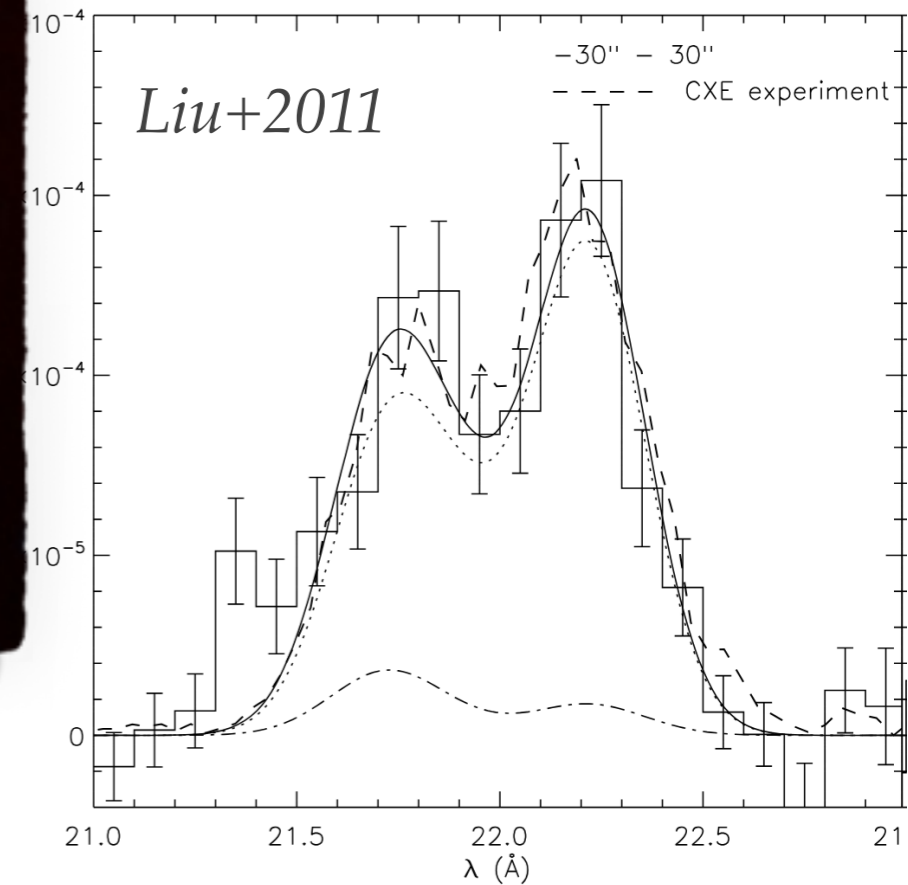
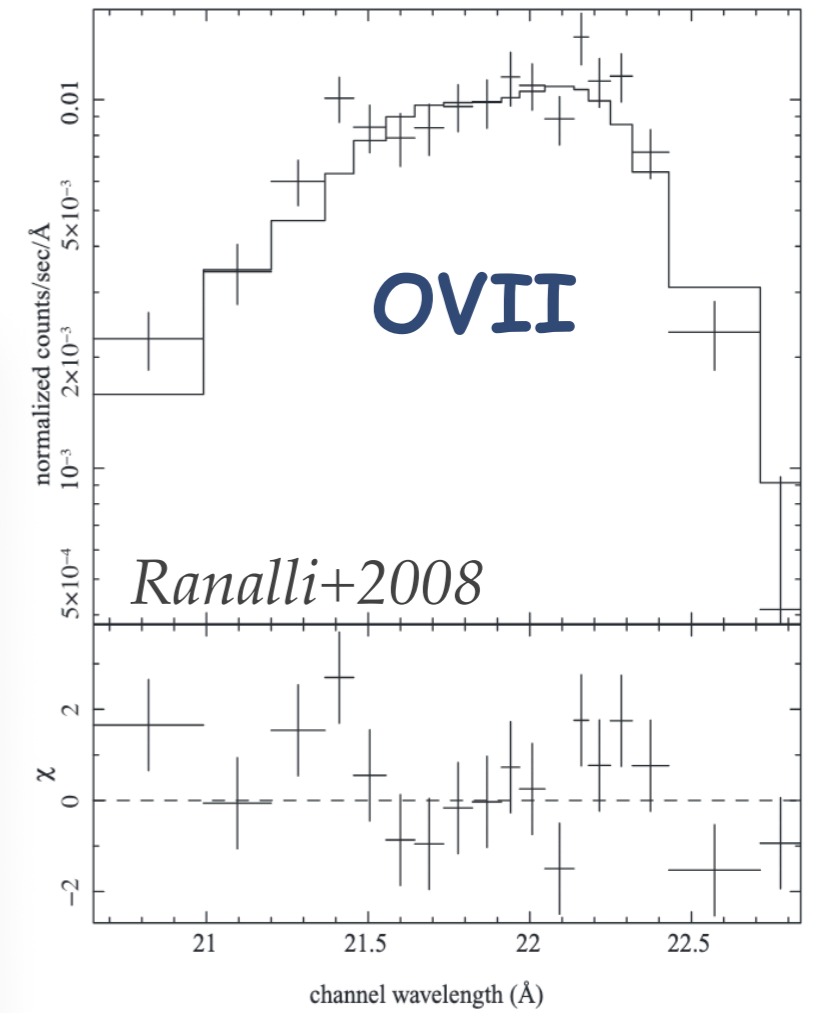
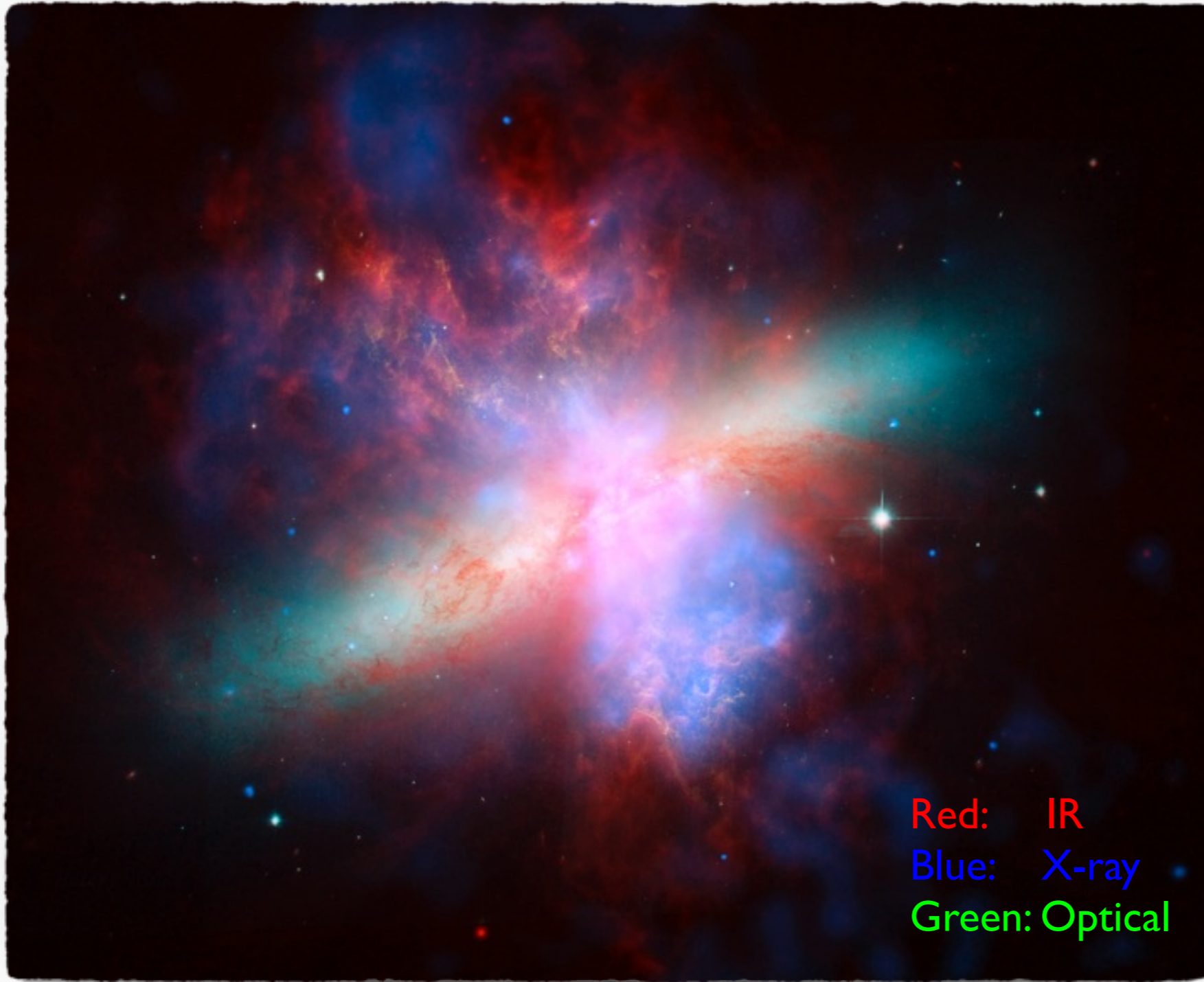
20 June 2024, Volos, Greece

**Trace the interplay between Hot and cold gas**

**—> An important coolant**

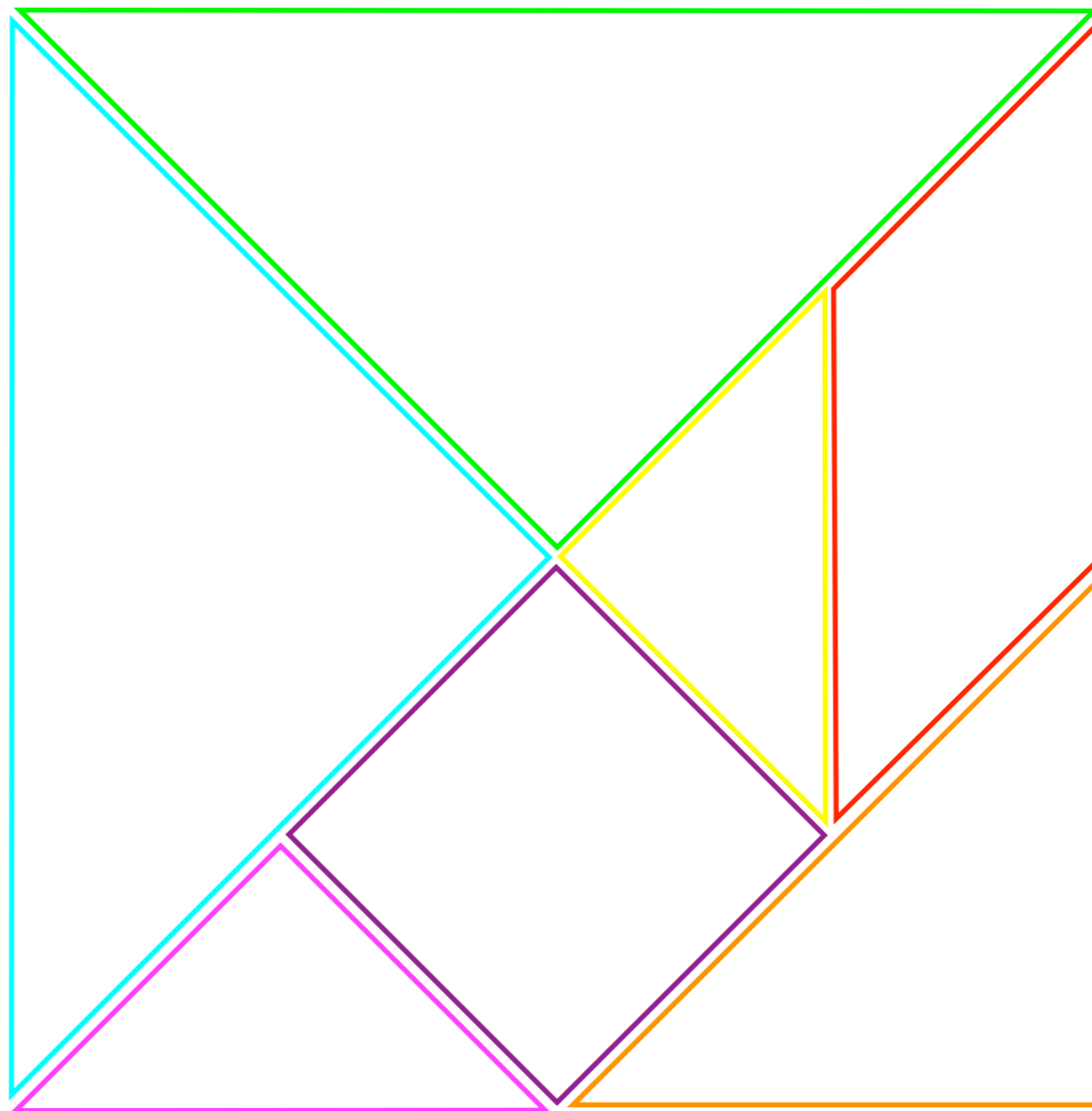
**—> Correct hot gas properties**

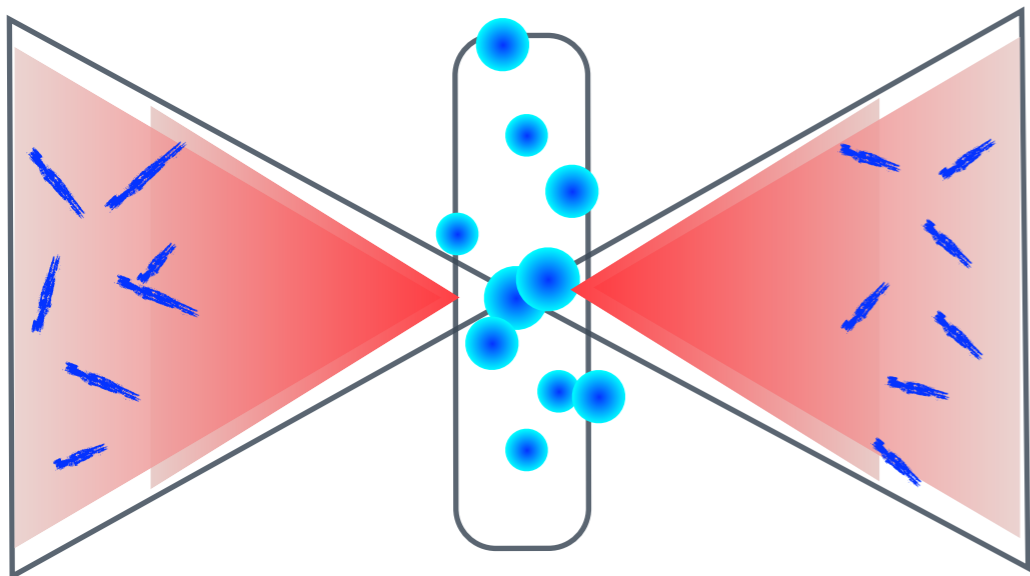
# Galactic superwind



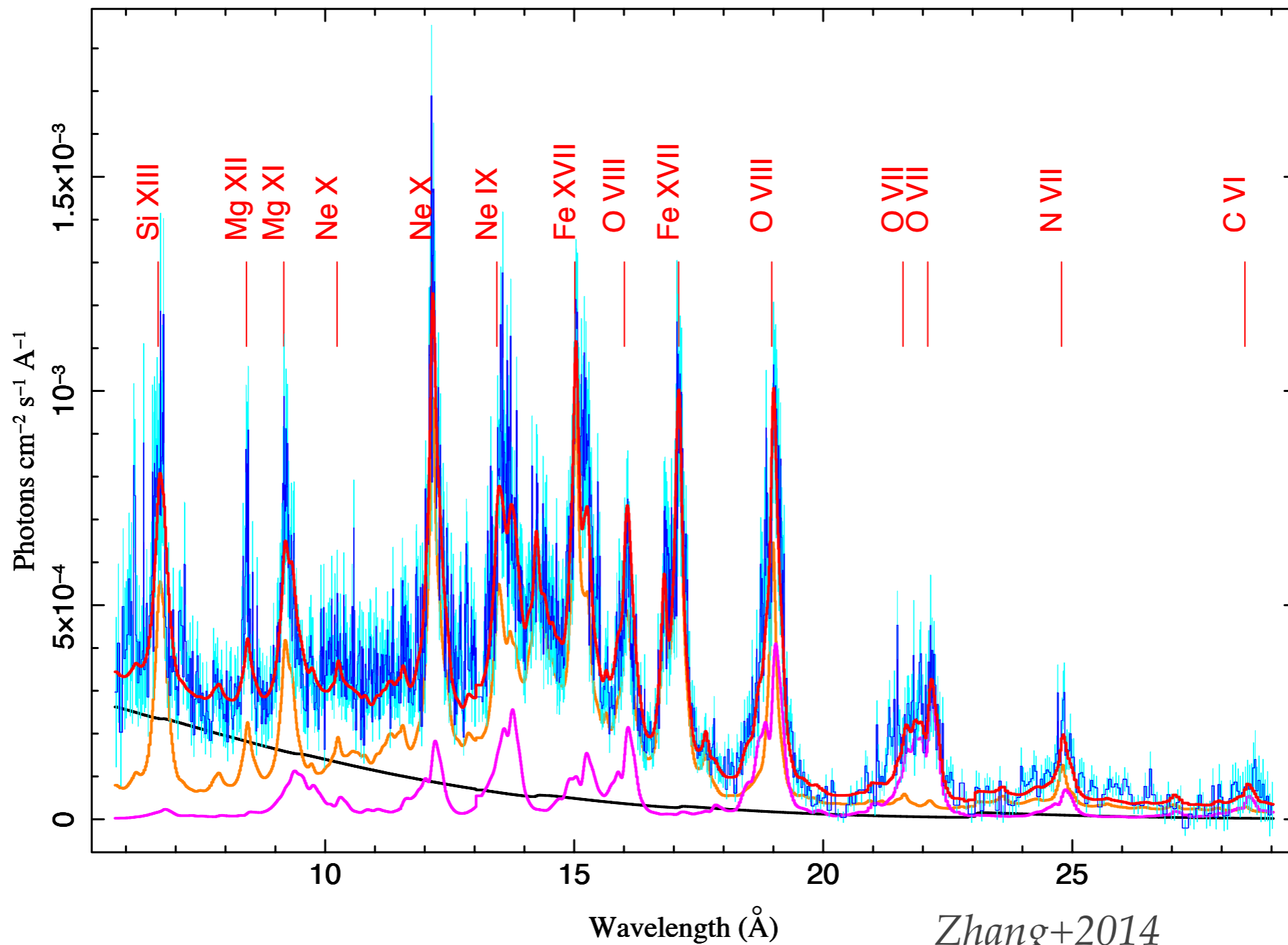
*Credit: X-ray: D.Strickland; Optical: The Hubble Heritage Team; IR: C. Engelbracht*

# Tangram





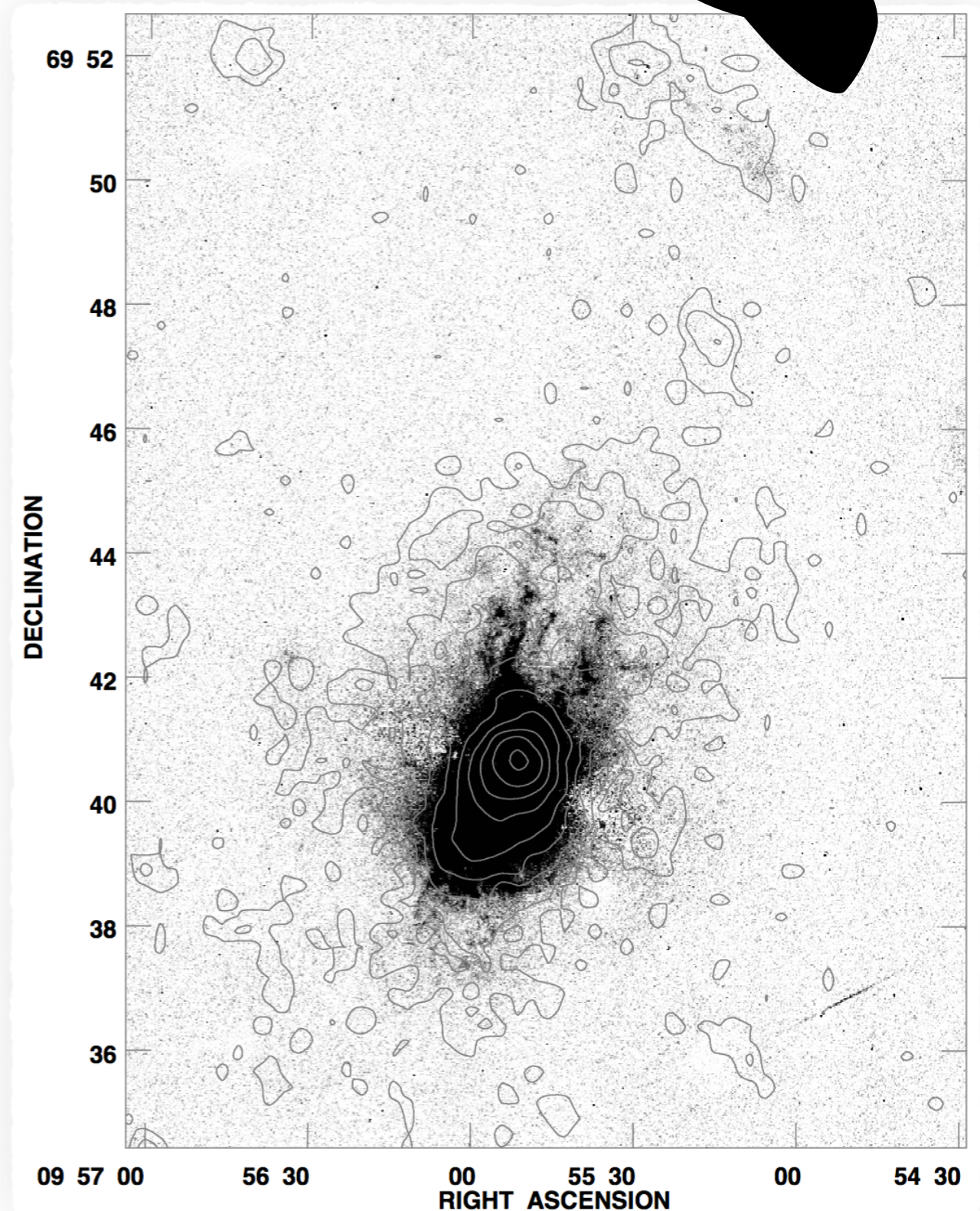
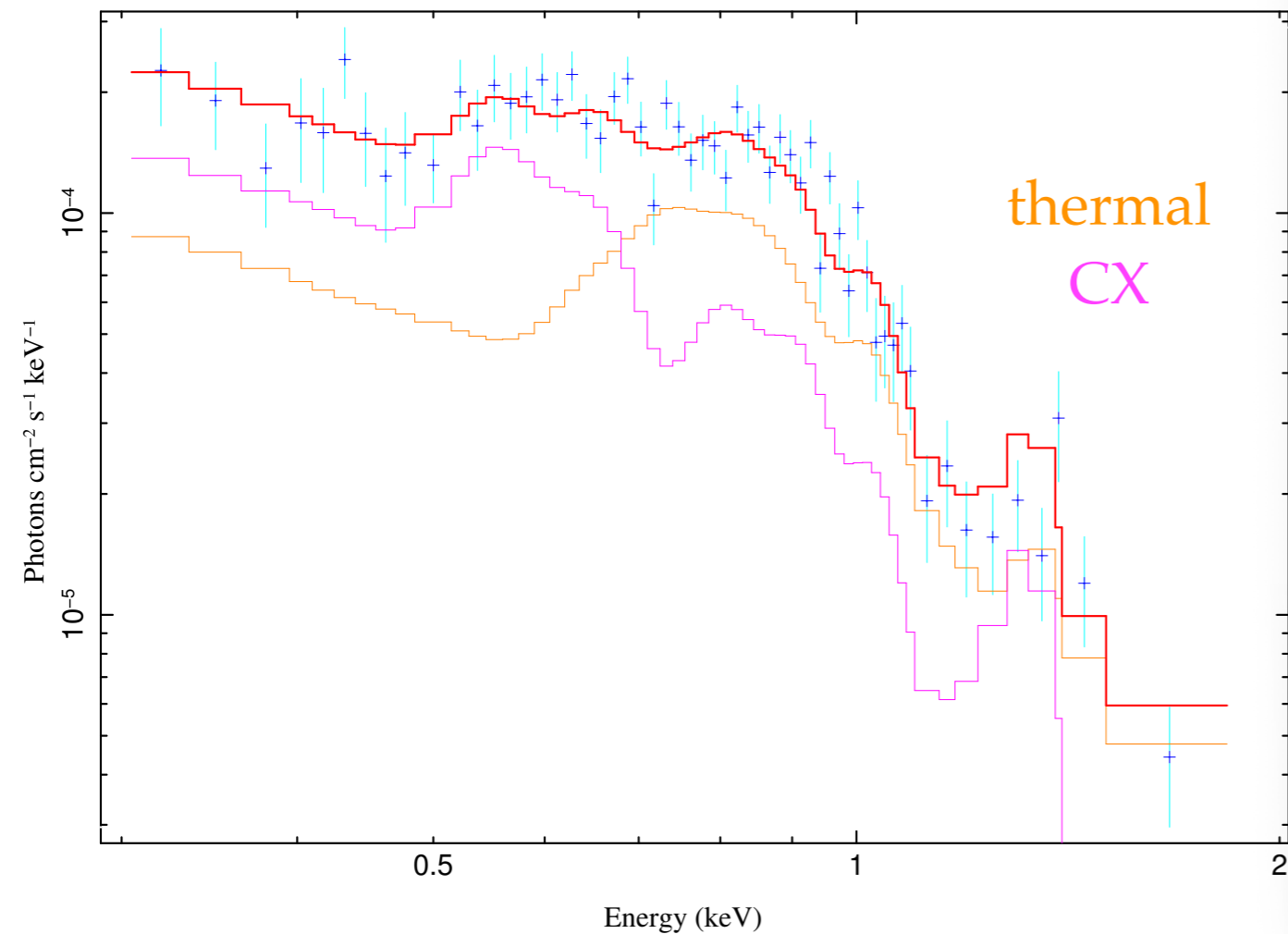
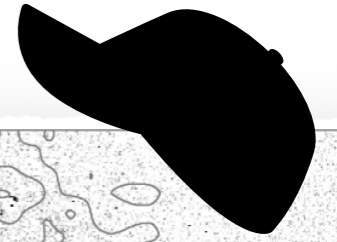
$$\text{abs1} * \text{pow} + \text{abs2} * (\text{thermal} + \text{CX})$$



0.6 keV + 0.2 keV (Tsuru+2007)

0.6 keV (50%) + CX (50%)

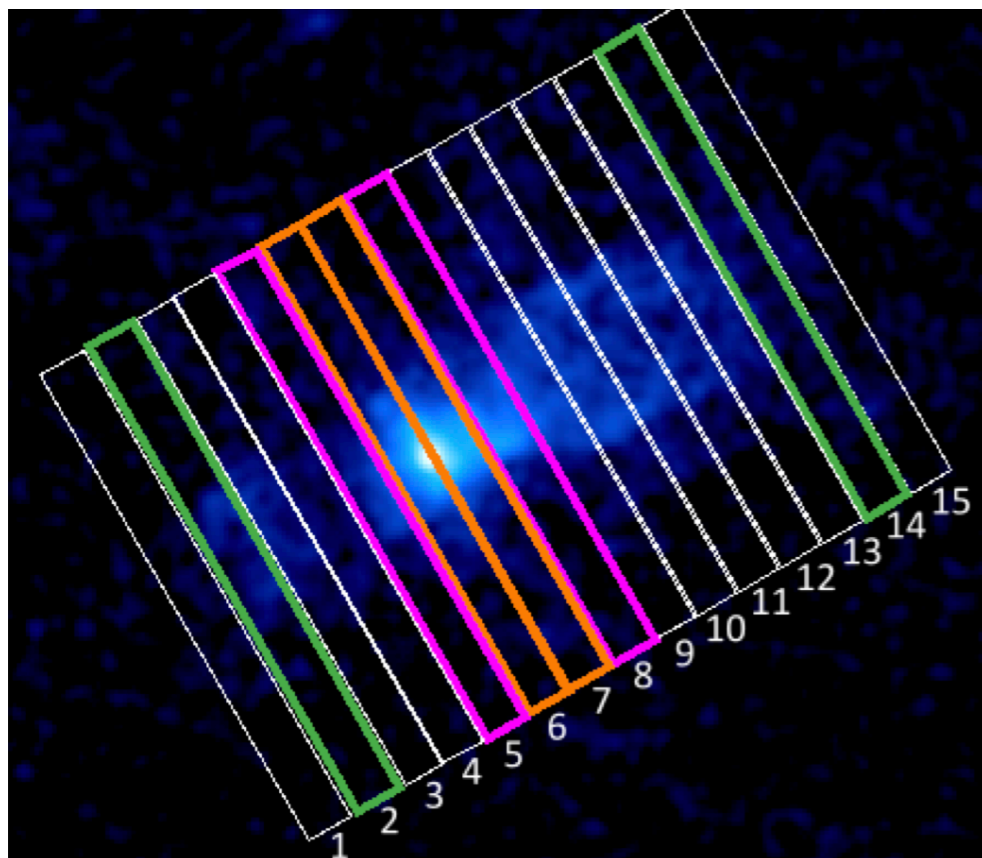
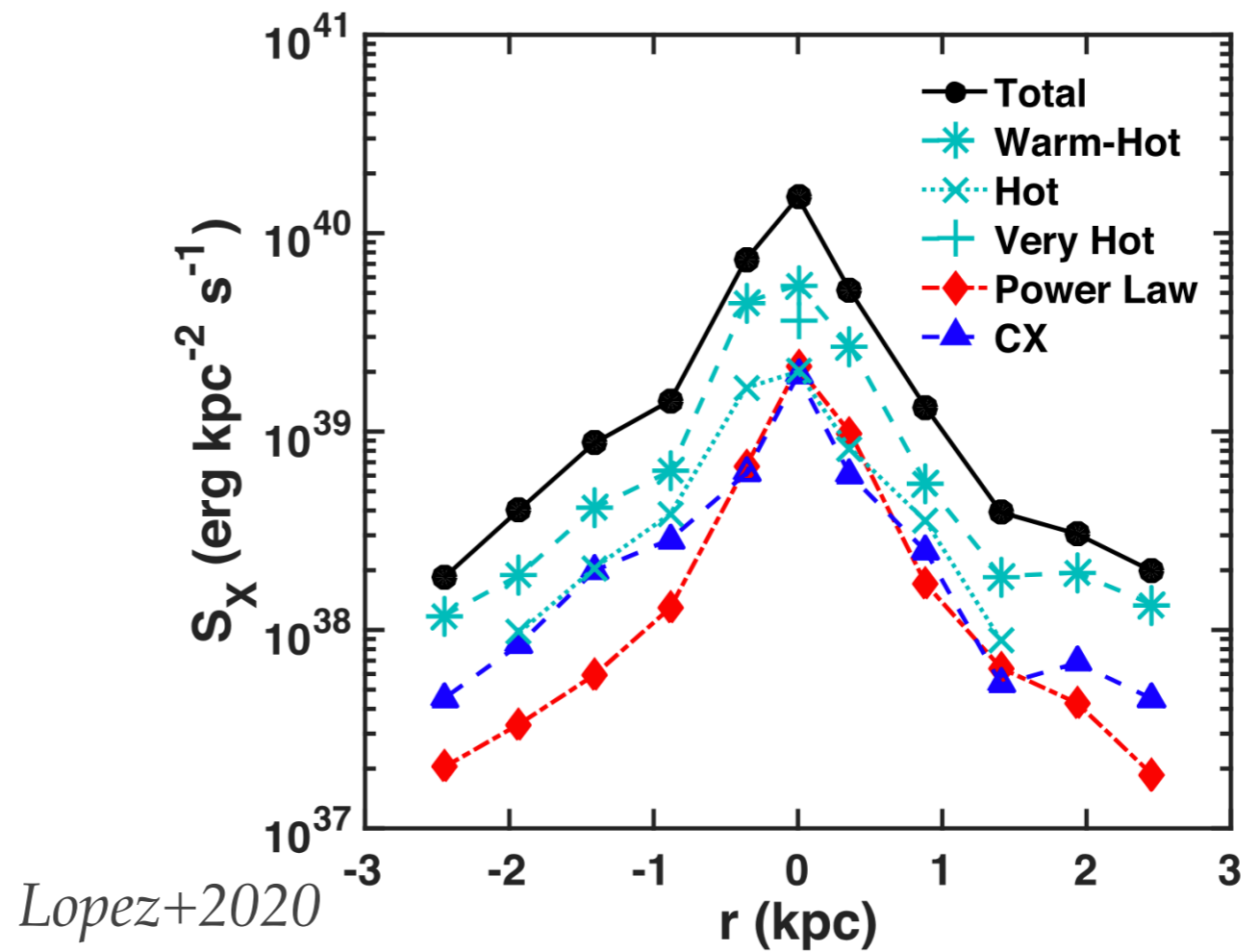
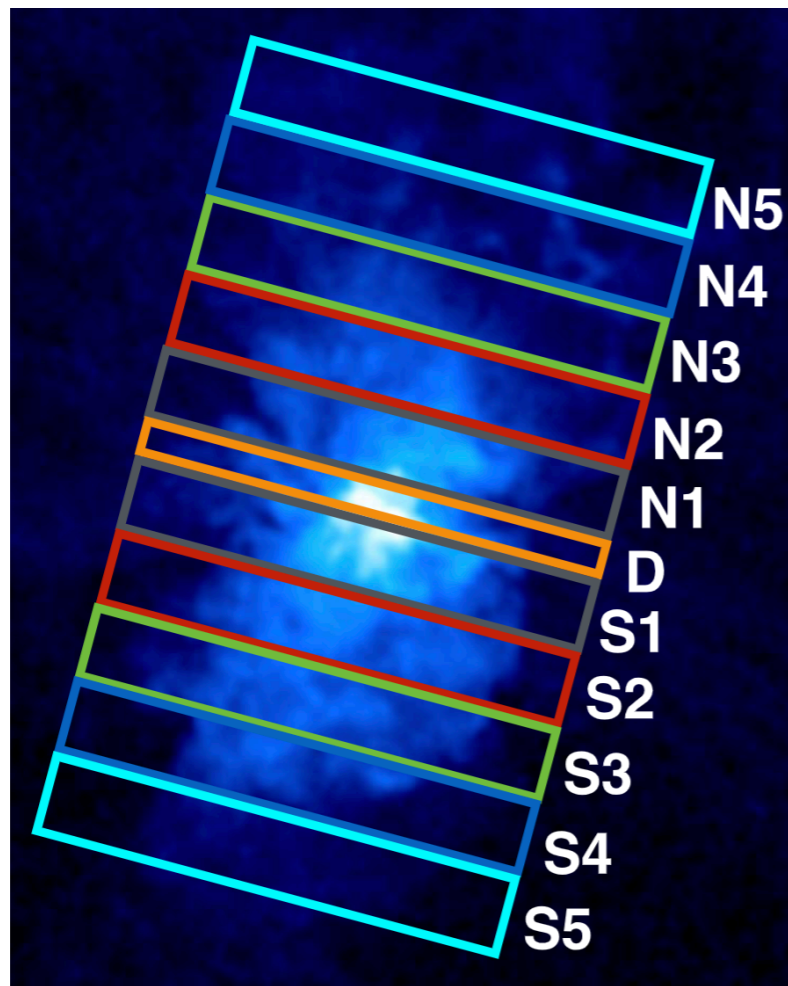
**Wind went into Cap region**



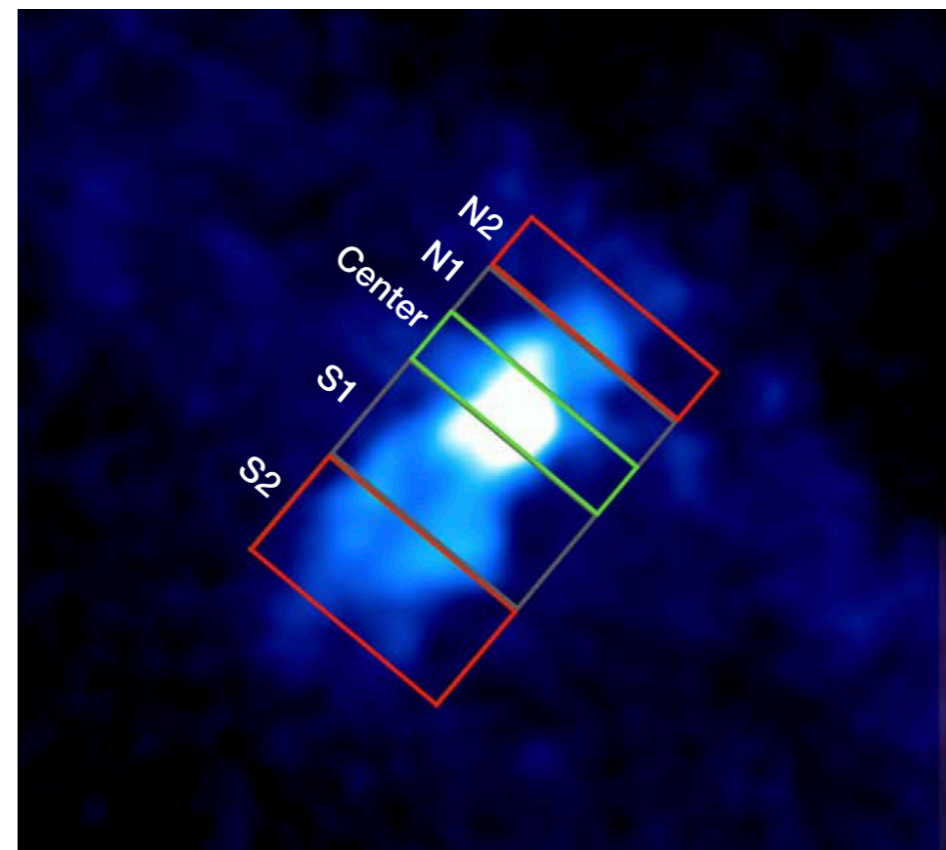
*Lehnert+2009*



kT &  
Metal

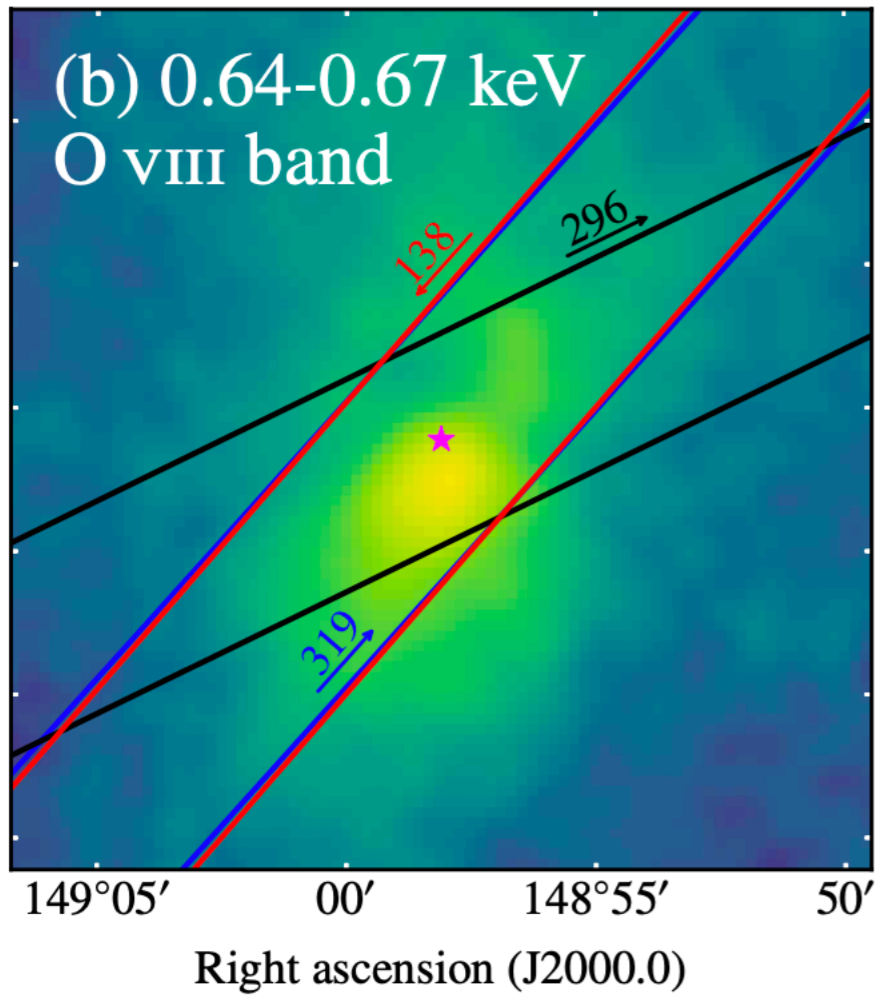


*Barrera+2024; NGC4945* CX 12%



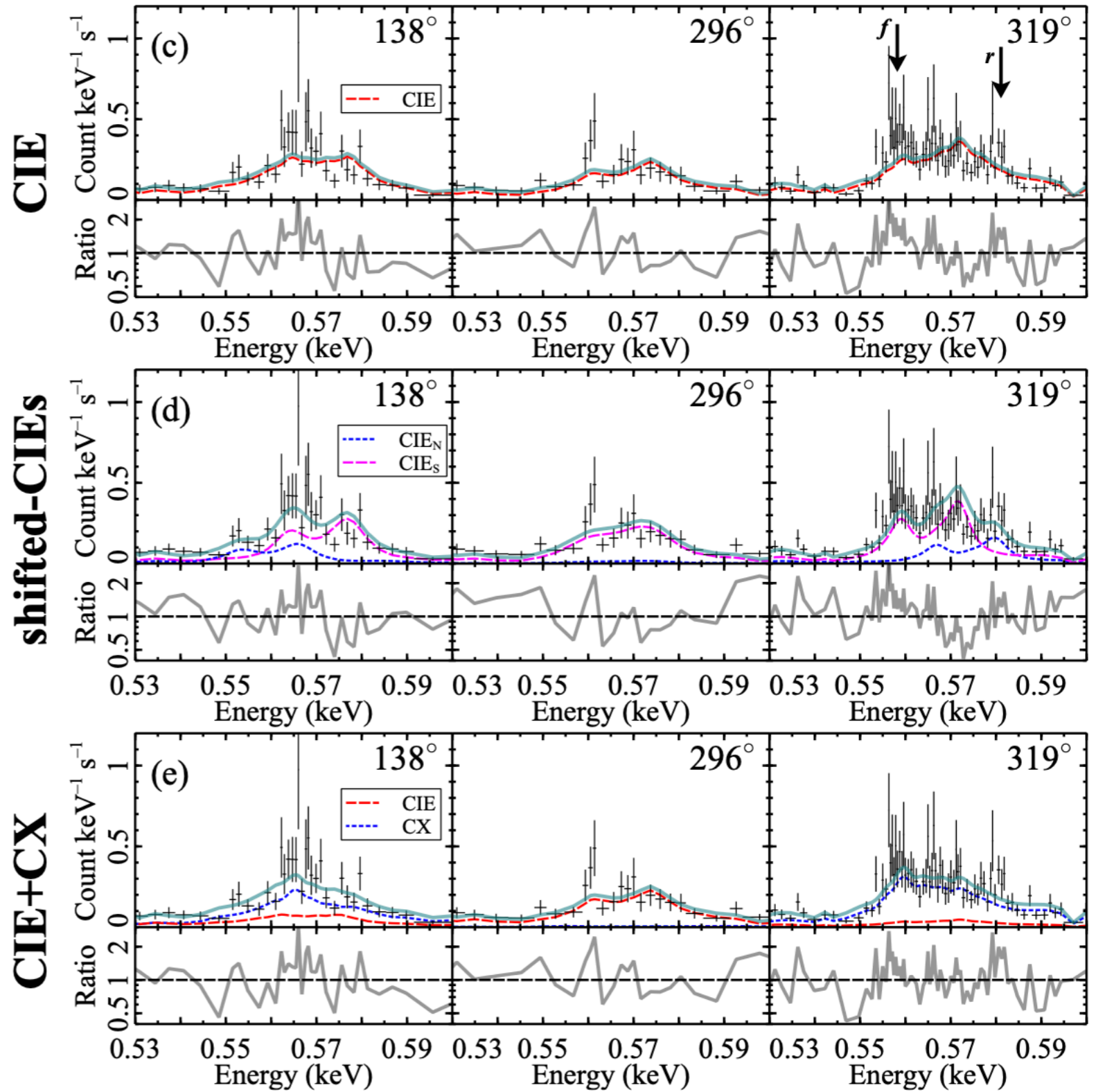
*Lopez+2023; NGC253* 20%–42%

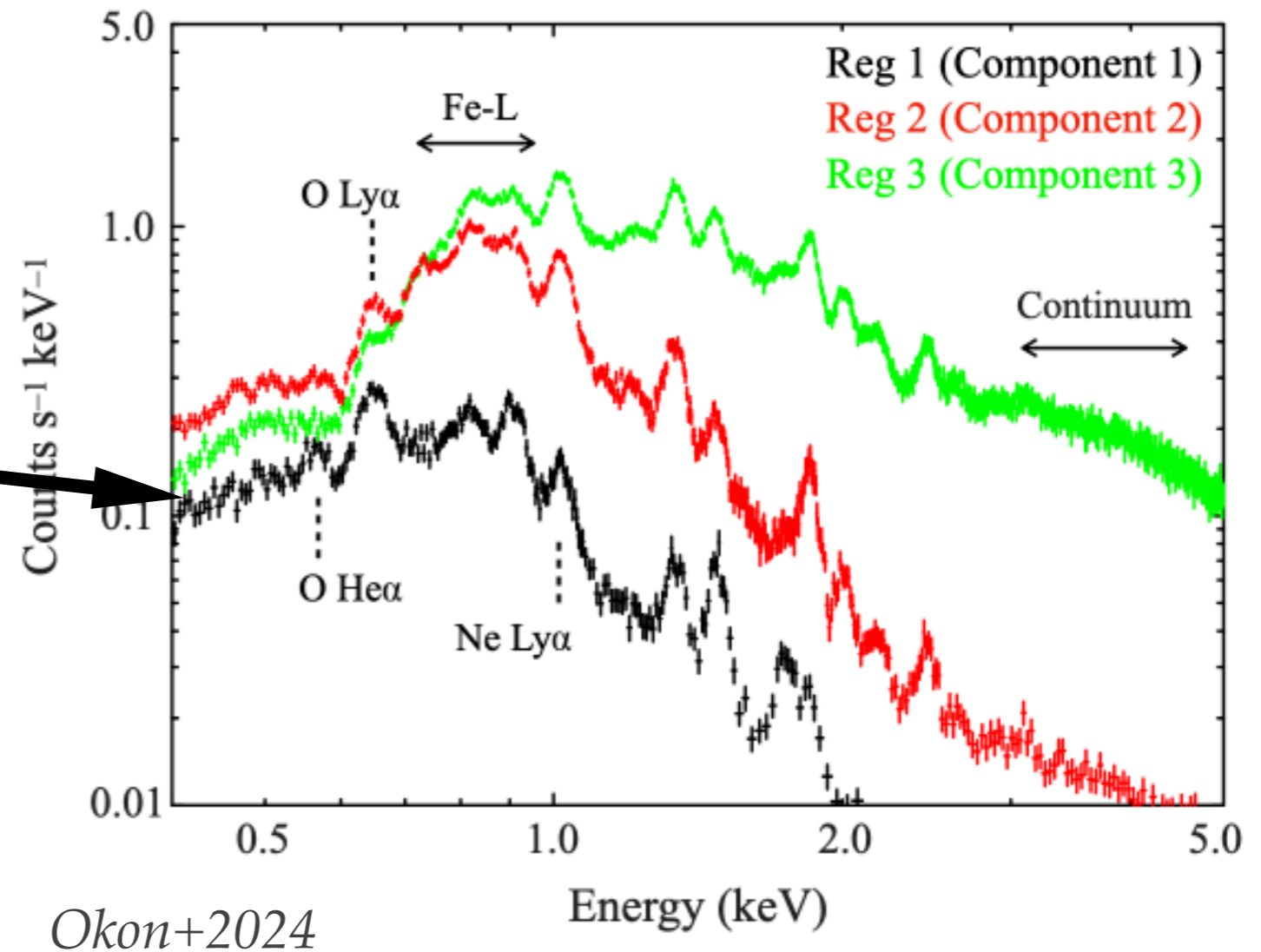
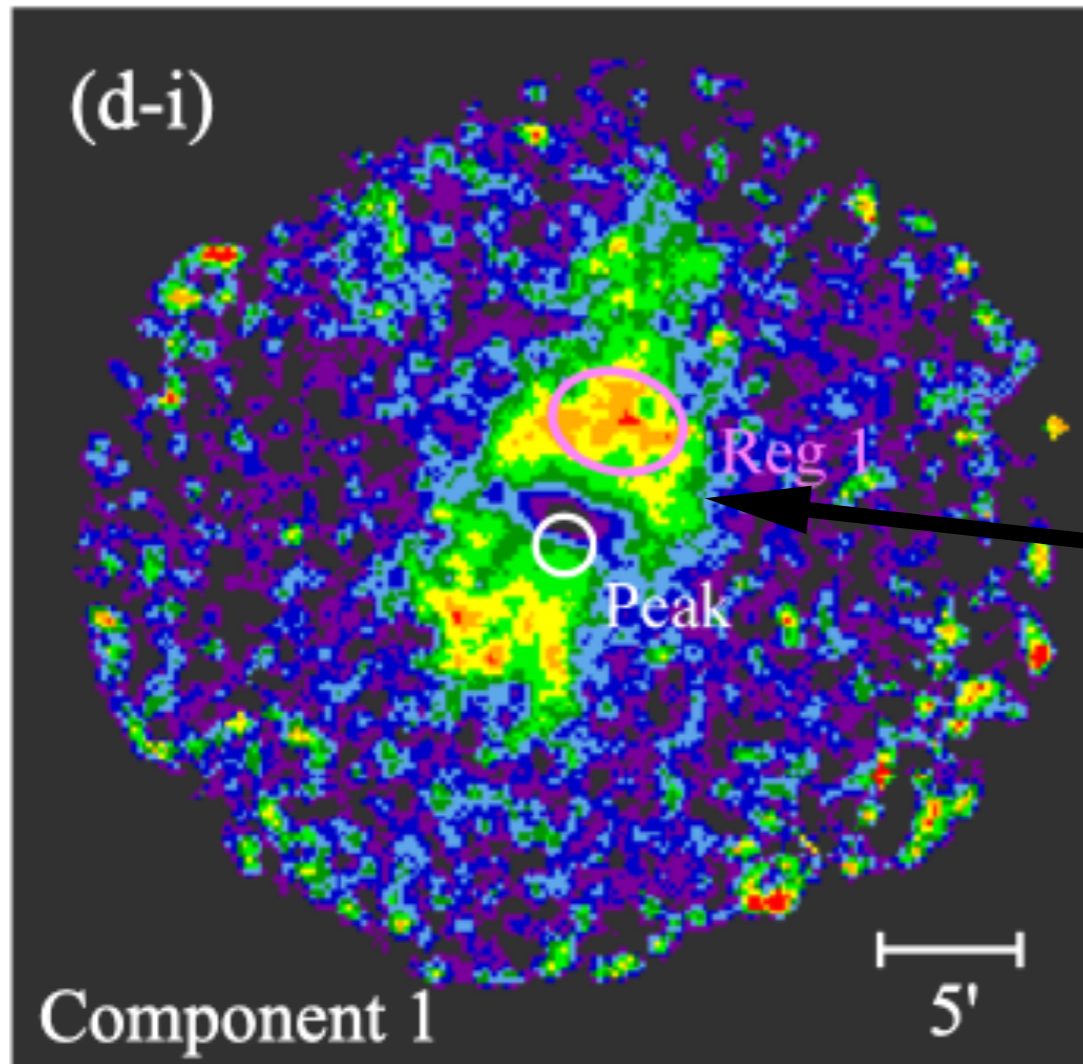




- Hard to see where is the CX
- The CX may be not true

## O VII triplet





- CX occurs in an extremely limited region near the interface zone.
- CX process mainly occurs in the layer whose temperature rapidly decreases due to thermal conduction ( $\sim 0.2$  keV).

0.3-1.1 keV

0.7-2.2 keV

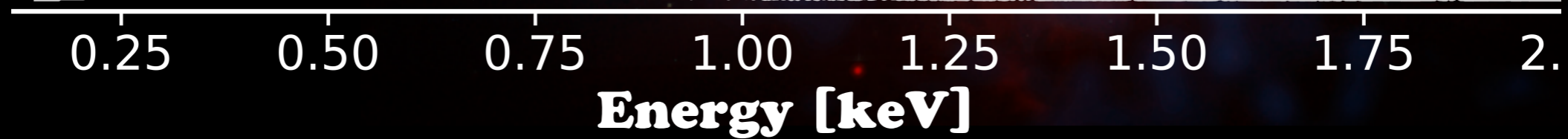
2.2-6 keV



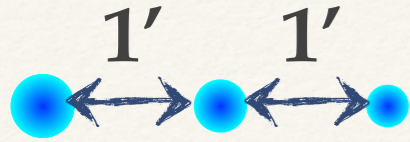
OVIII

OVII

GALEX UV

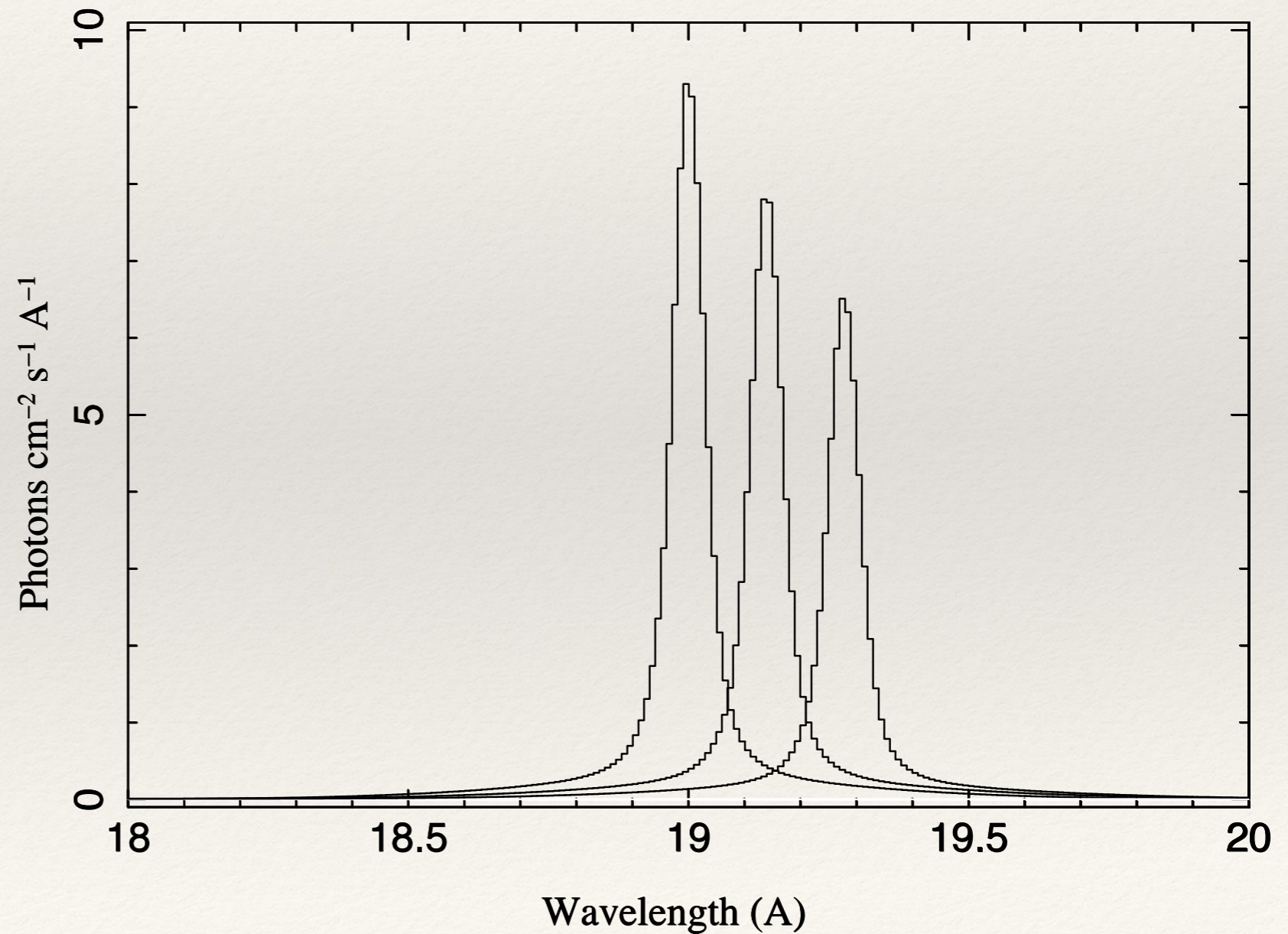
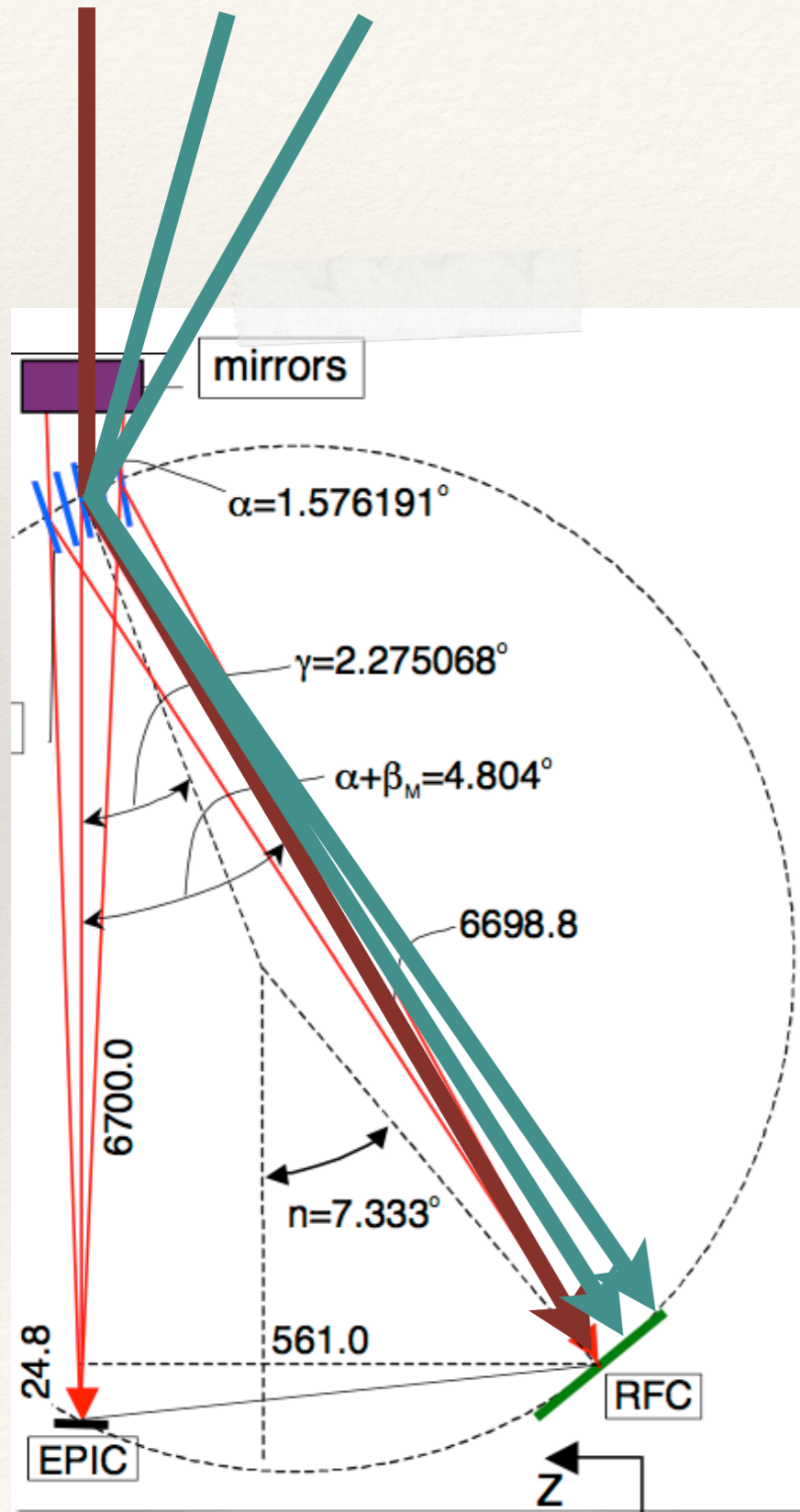


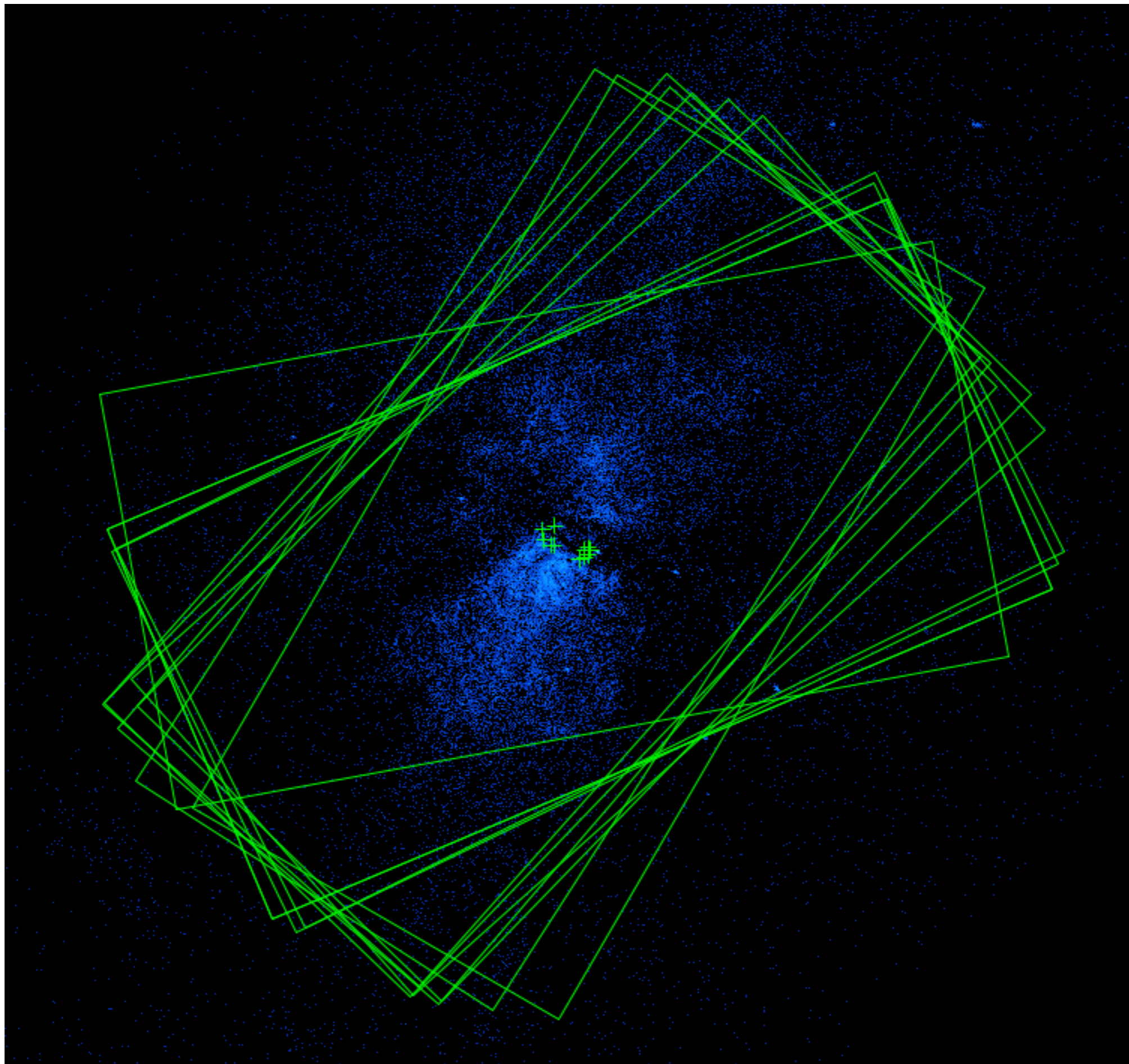
# Single Energy Photons



## Slit-less Grating records Spatial Info

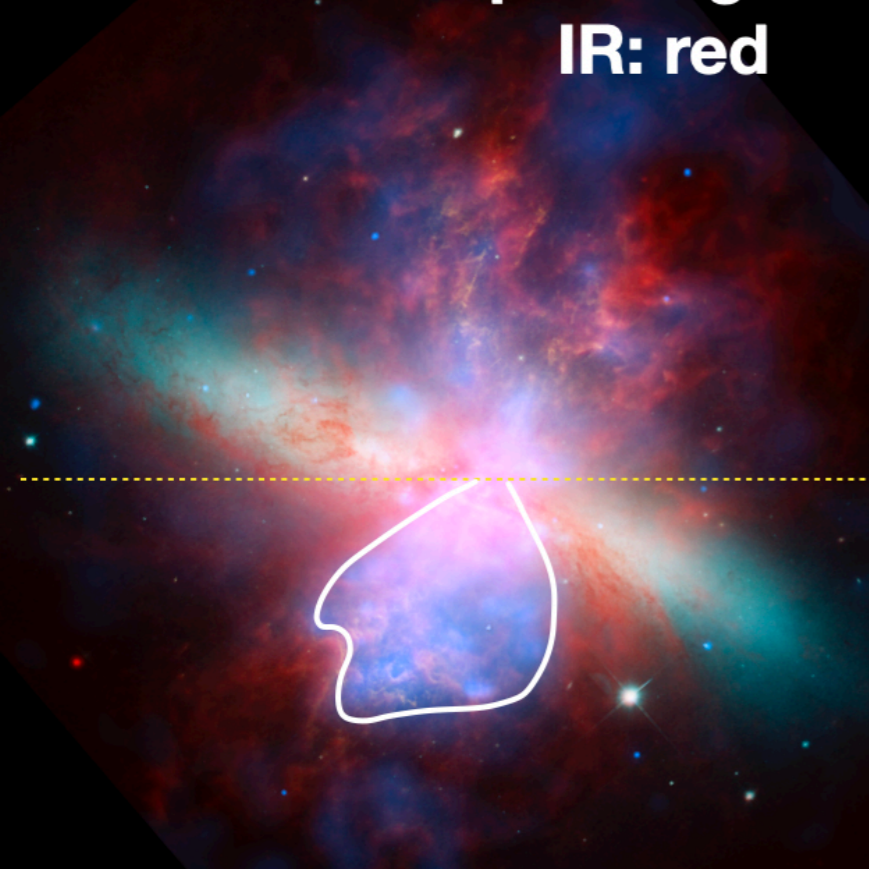
$$1 \text{ arcmin} \sim 0.138 \text{ \AA}$$





M82

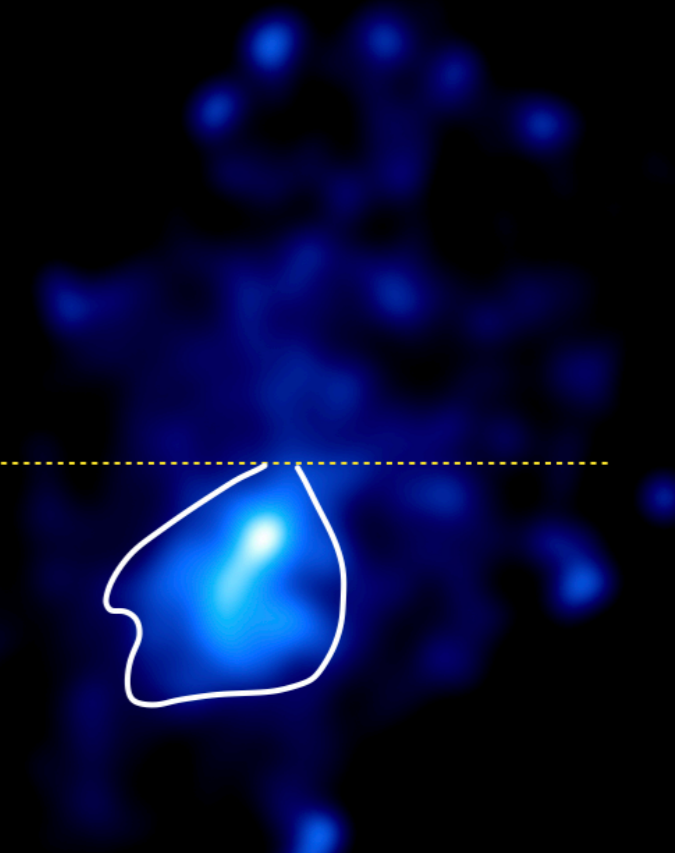
a) X-ray: blue  
Optical: green  
IR: red



b) XMM-Newton  
OVIII Ly $\alpha$  map

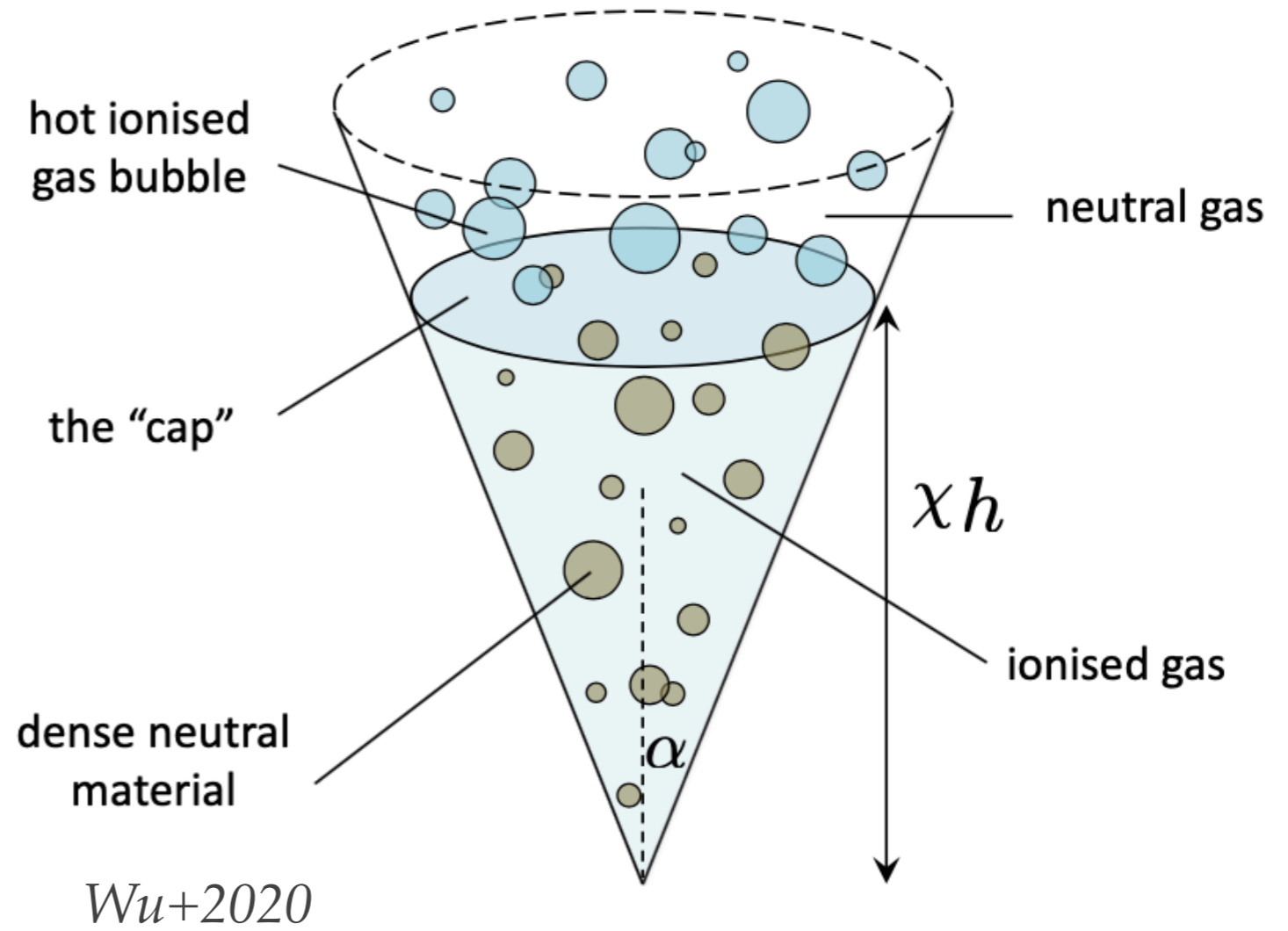


c) XMM-Newton  
OVII f line map



- Only data.
- OVII f is mainly in the southern wind, along the outflow axis.
- OVII f  $\rightarrow$  CX distribution.

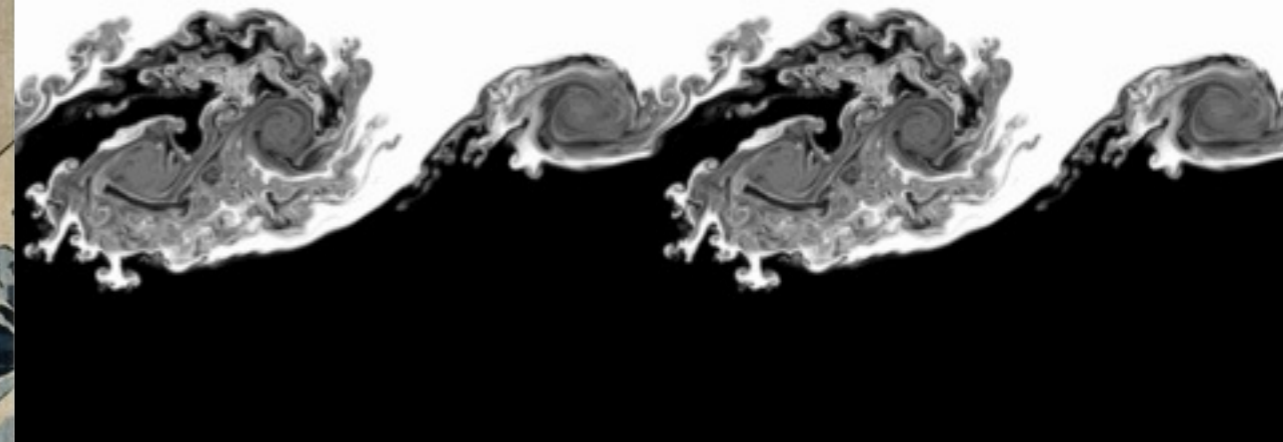
Where  
CX



- CX simply occurs at the surfaces of the cold clumps entrained and the cold shell.



# Turbulences at the interface



**for  $10^{15}$  cm scale**

**Ionization:**

$10^8$  s

**Recombination:**

$10^{12}$  s

**CX:**

a half year

**Conduction:**

a half year



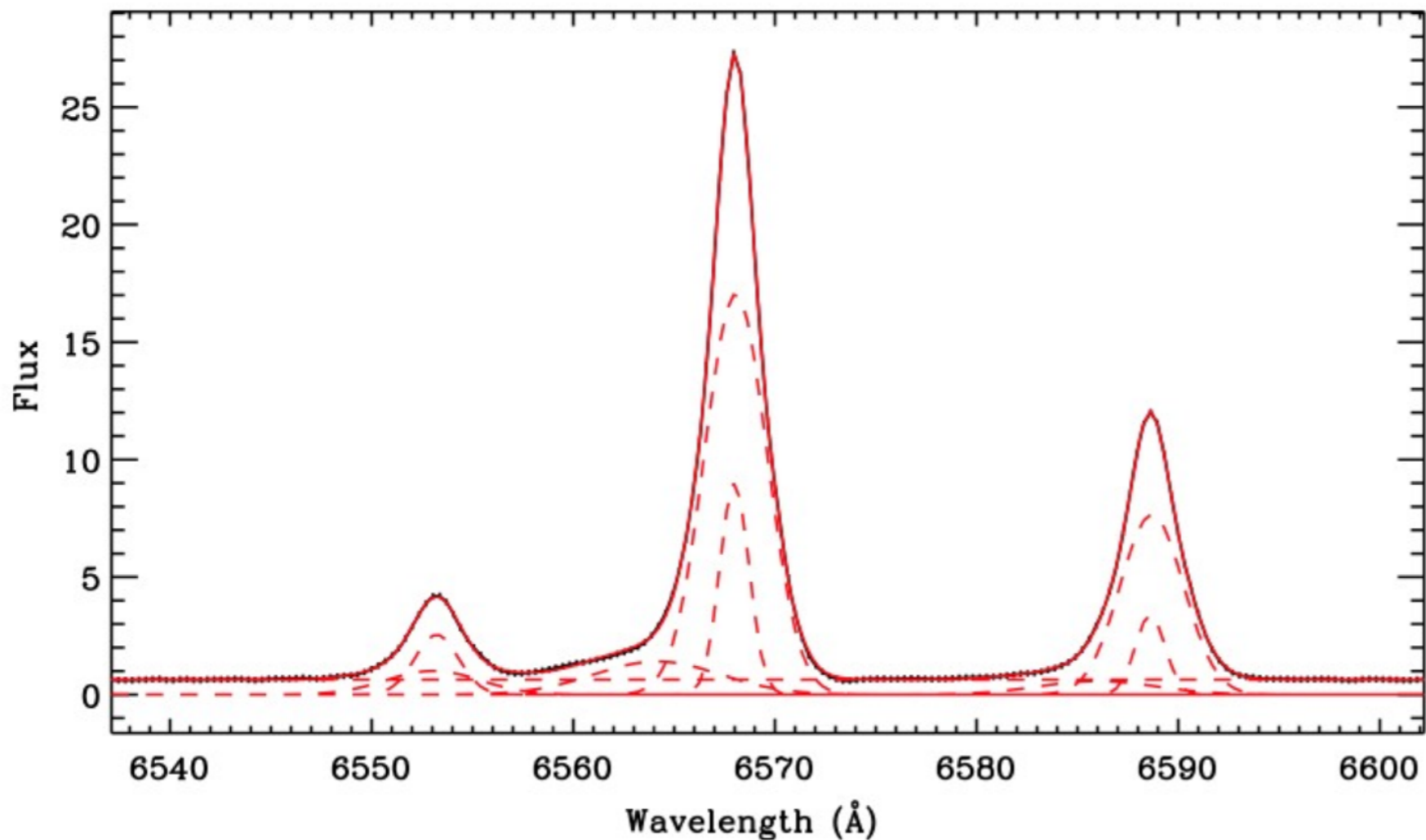
How

# Broad H $\alpha$ component

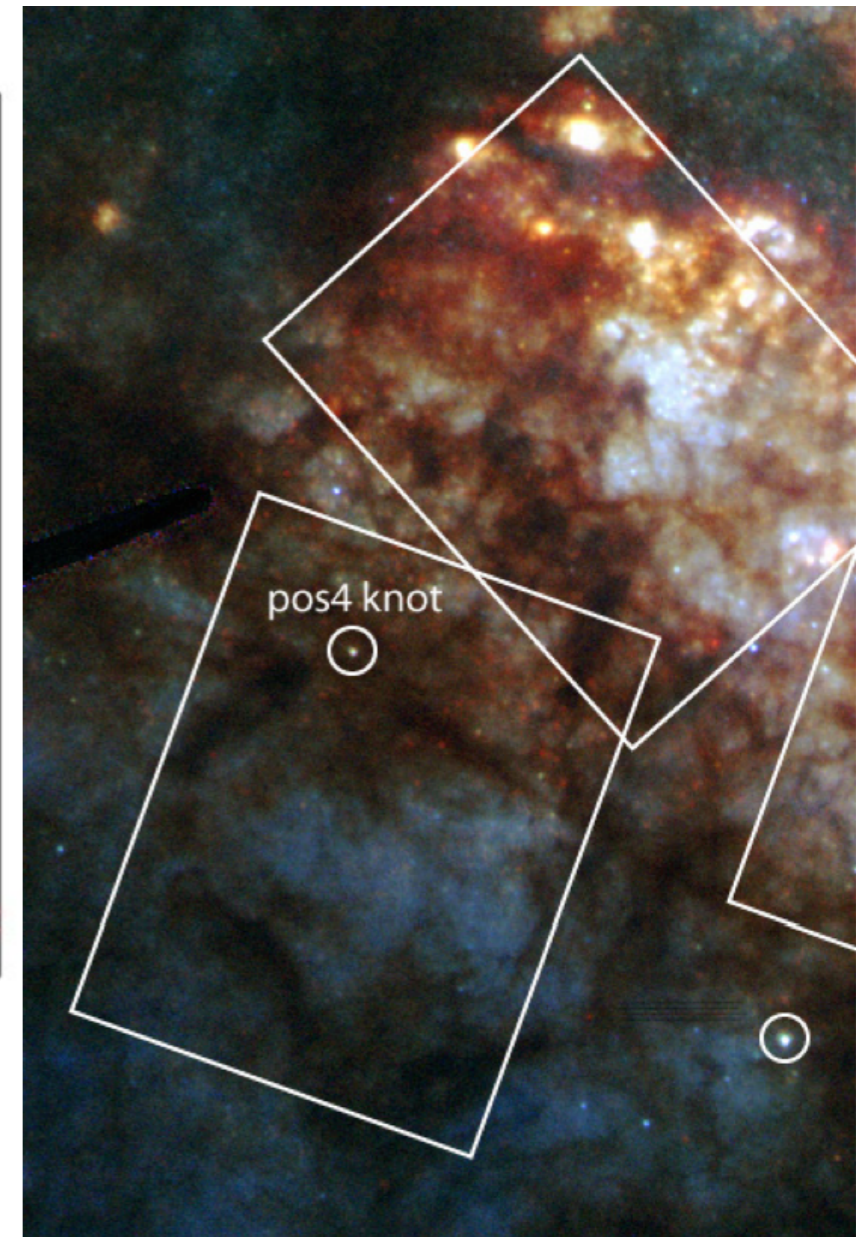
Component	H $\alpha$ FWHM (km s $^{-1}$ )	H $\alpha$ Radial Velocity (km s $^{-1}$ )
C1	156	40
C2	333	-138
C3	53	37

CX contributes **3%** H $\alpha$

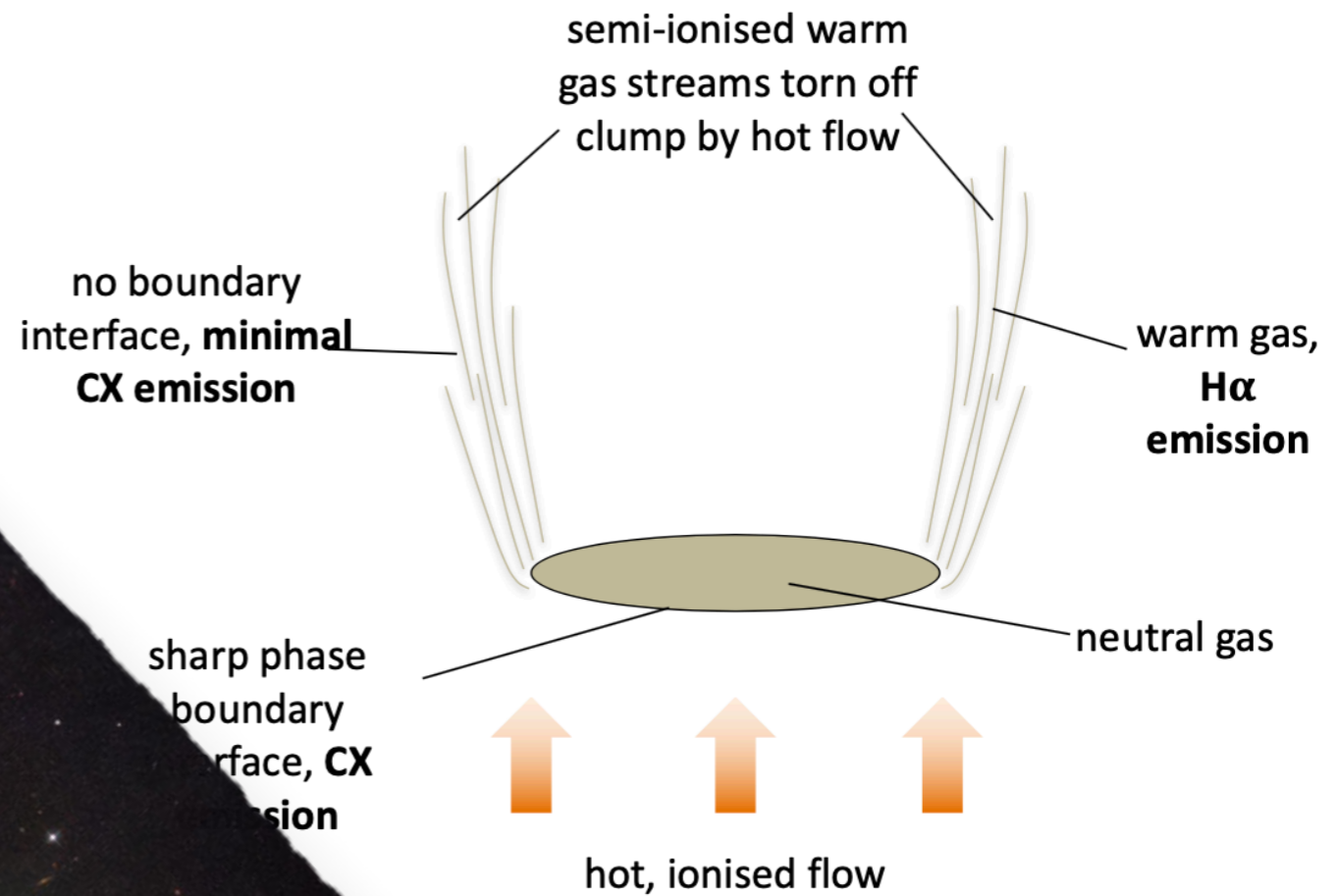
Southern wind



Westmoquette et al. 2009: Gemini GMOS-IFU



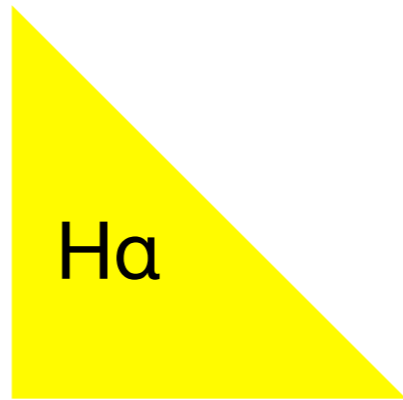
# H $\alpha$ Not related to CX directly



Wu+2020

Credit: NASA & ESA

Acknowledgement: M. Mountain (STScI)  
Image Credit: NASA, ESA



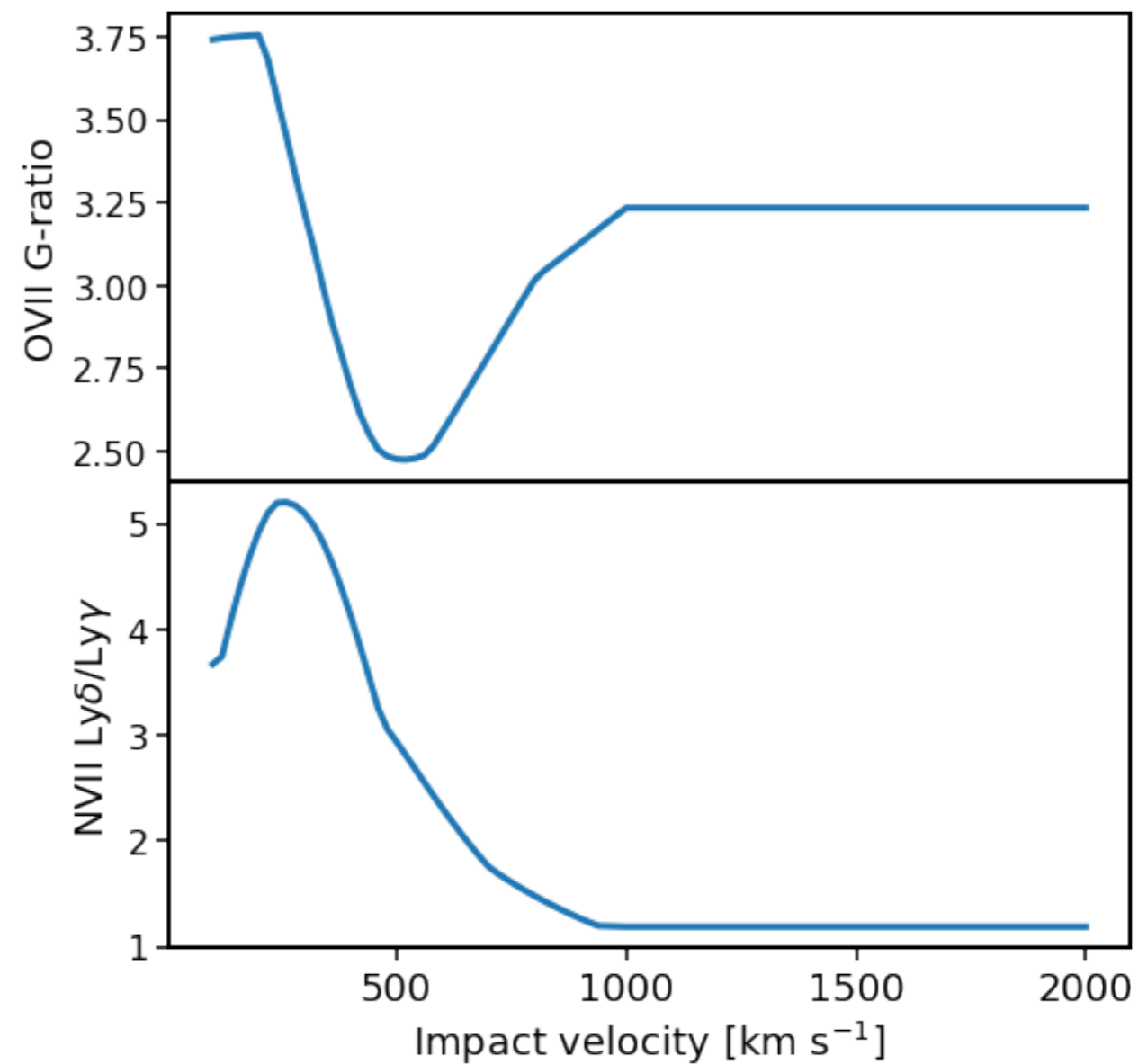
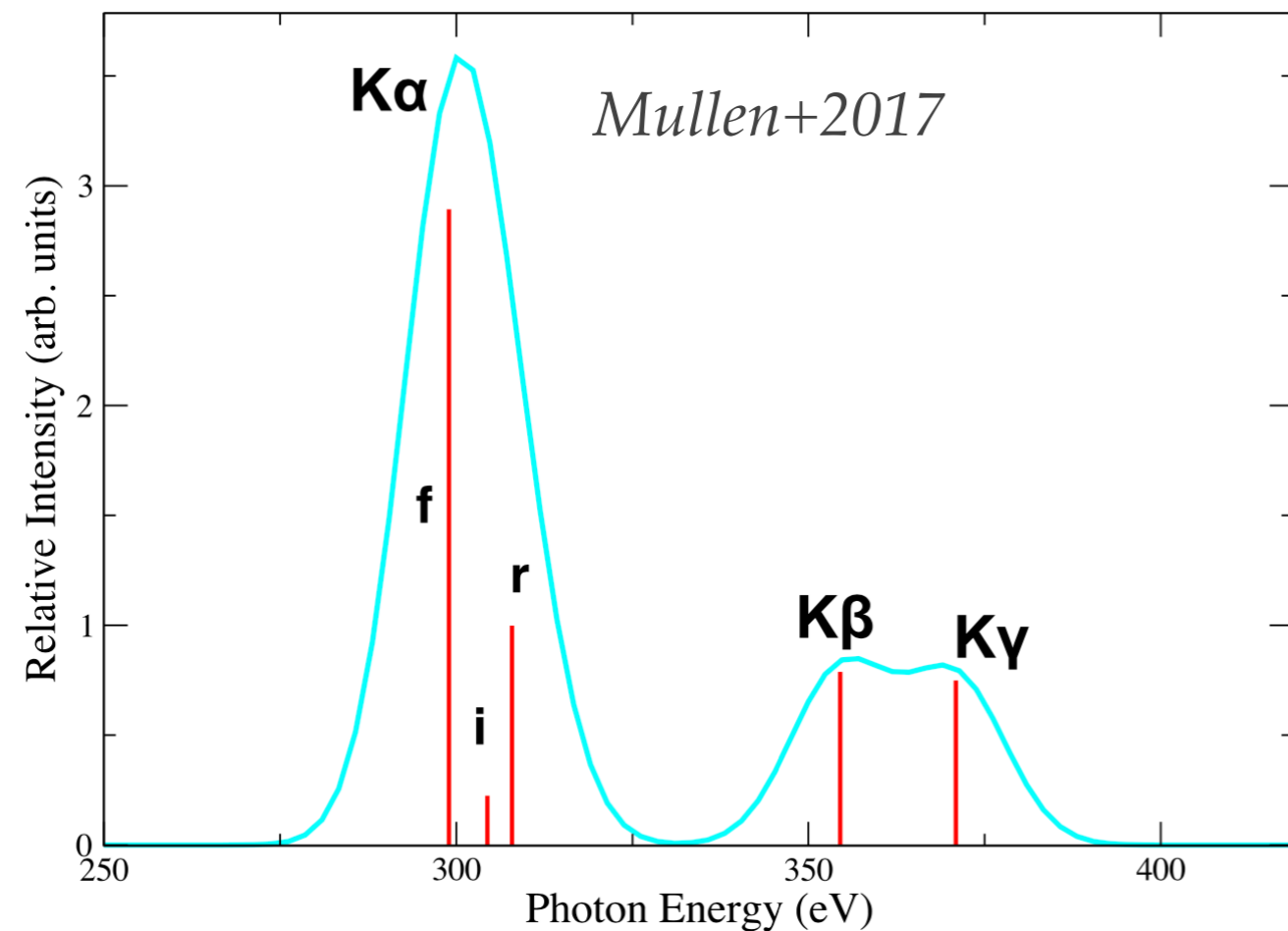
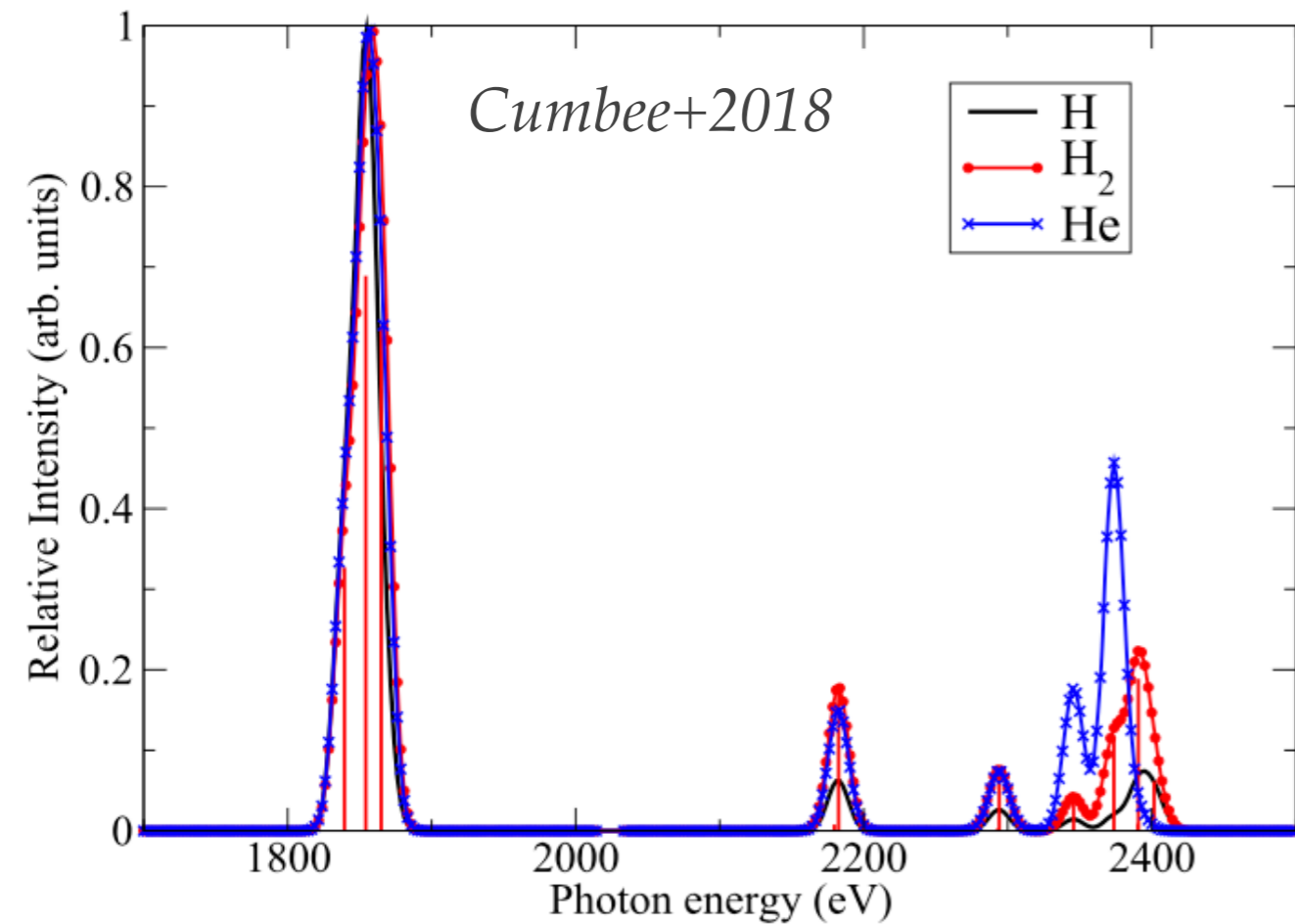
Ha

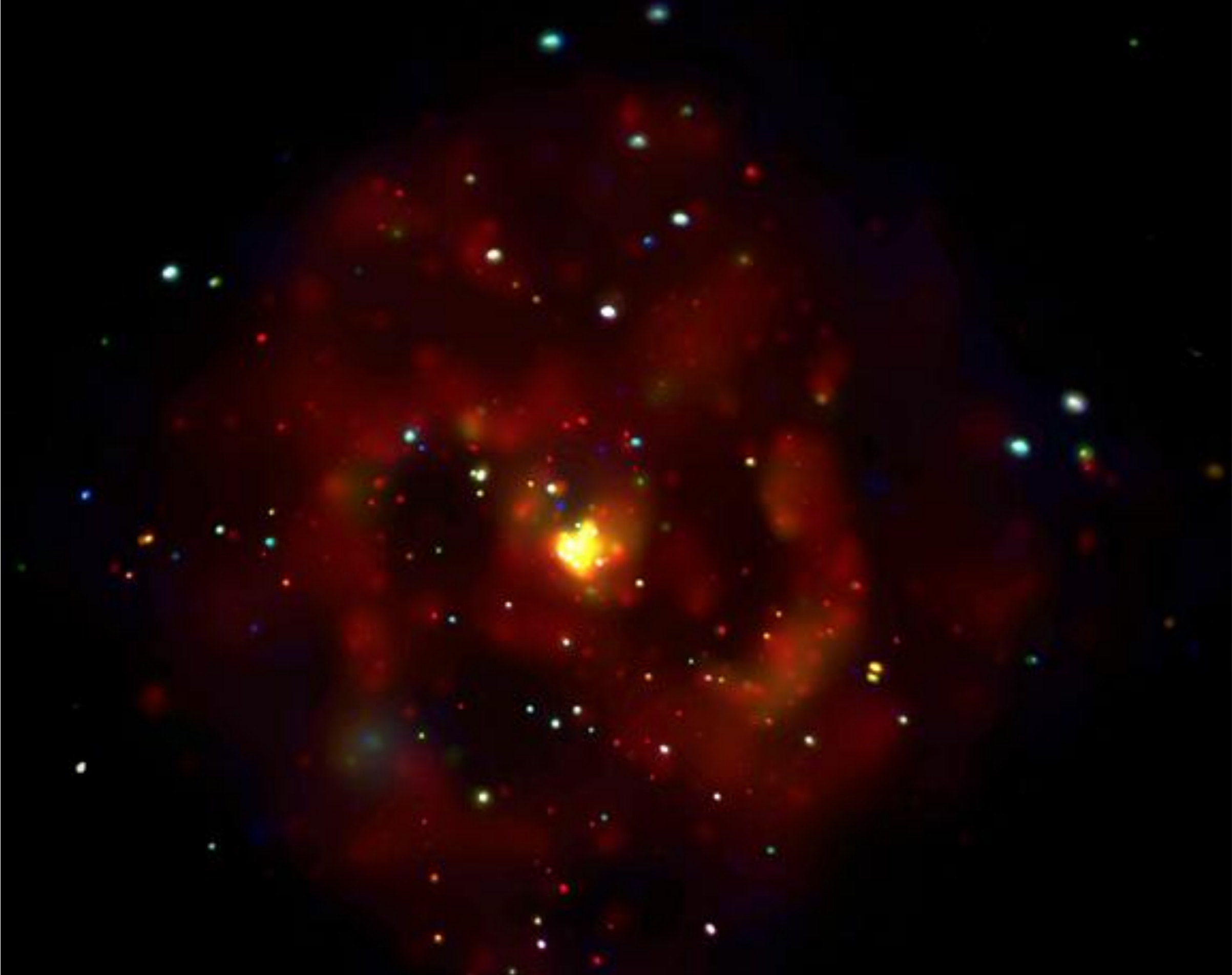
Assuming interacting velocity: 500 km/s

The interface area of hot & cold gas ( $\sim 10^{45} \text{ cm}^2$ )  
is one order of magnitude larger than  
the superwind geometric area ( $\sim 10^{44} \text{ cm}^2$ )



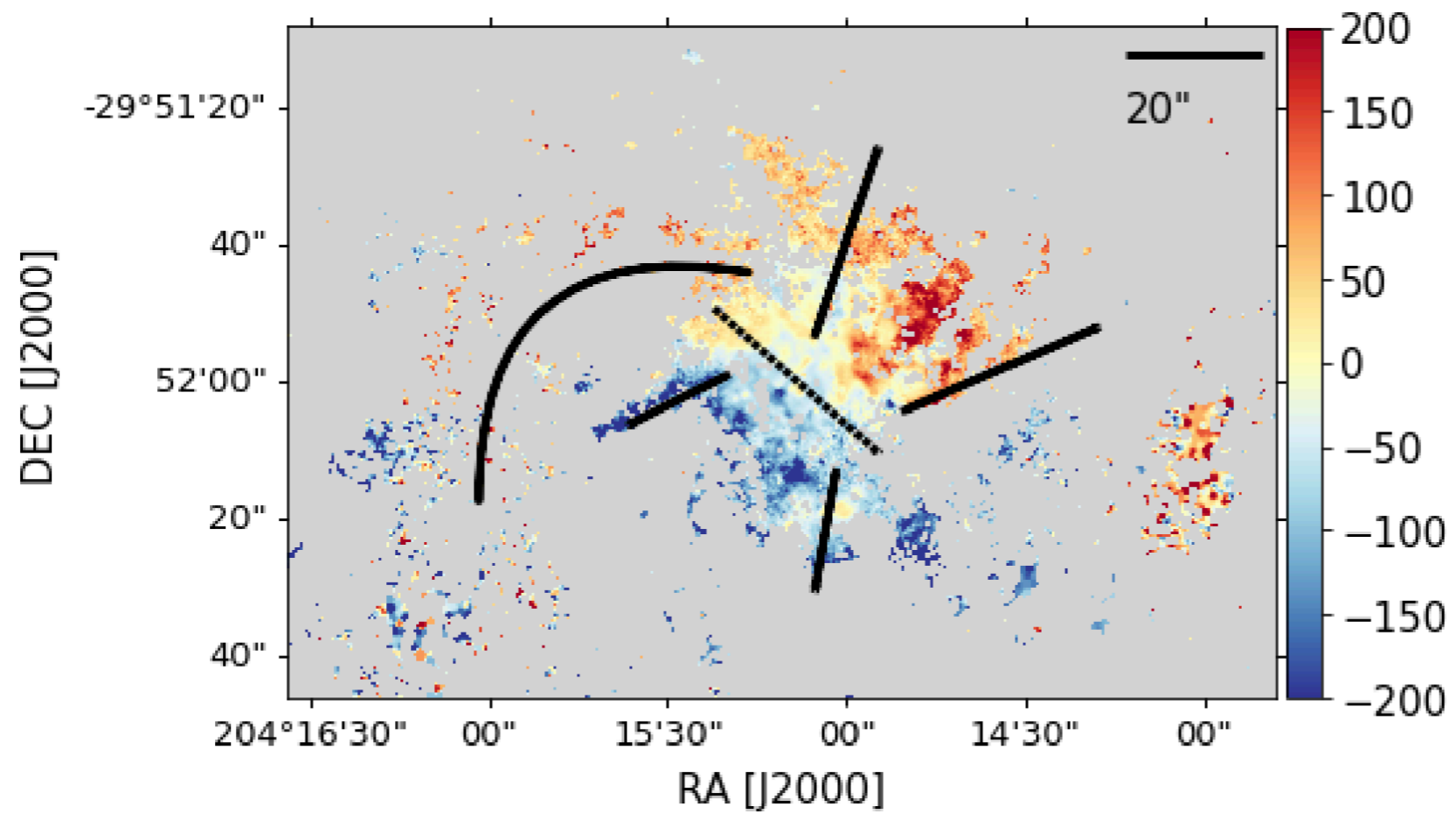
# Constrain velocity using CX



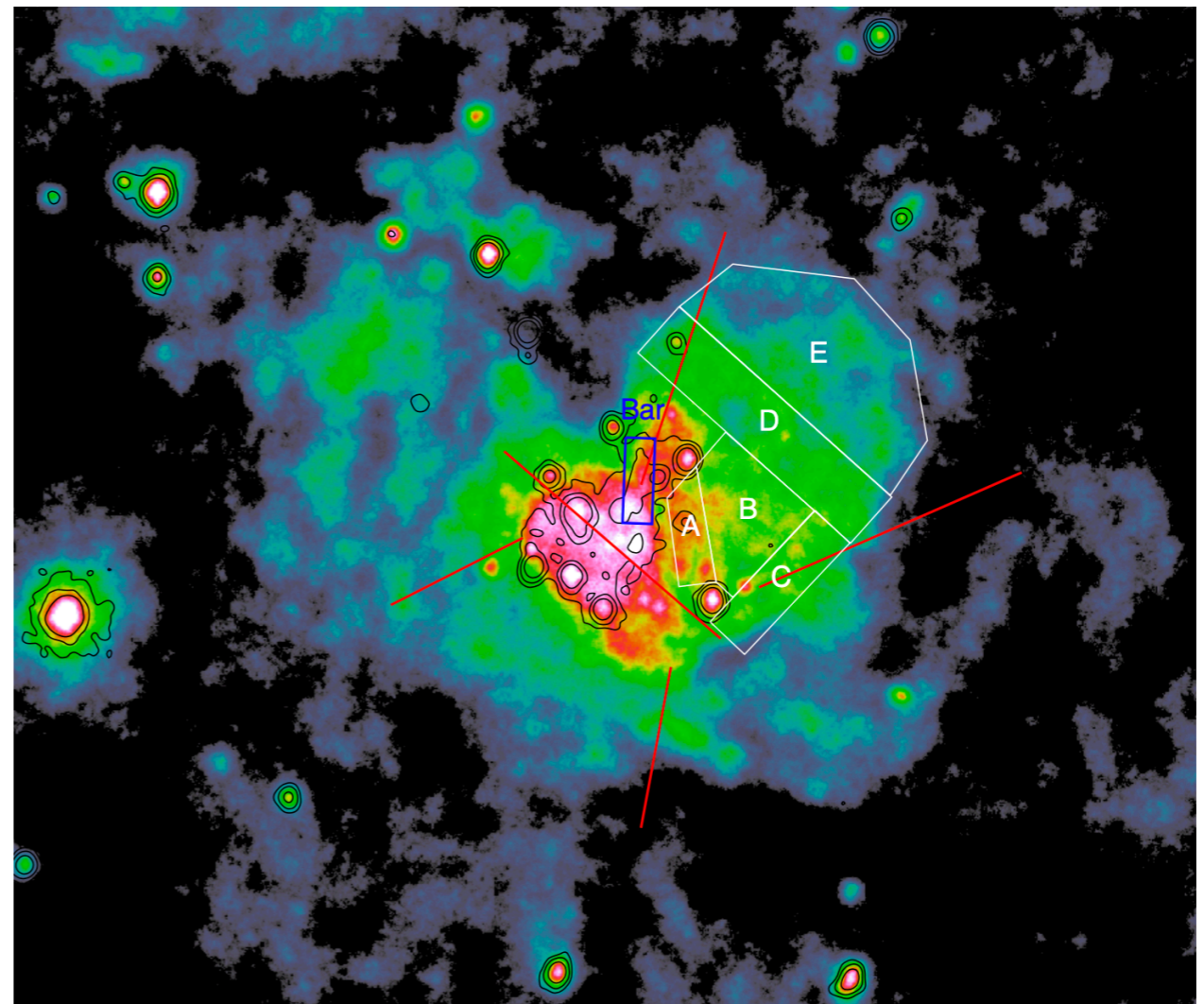


*Credit for above images: NASA/CXC/U.Leicester/U.London/R.Soria & K.Wu*



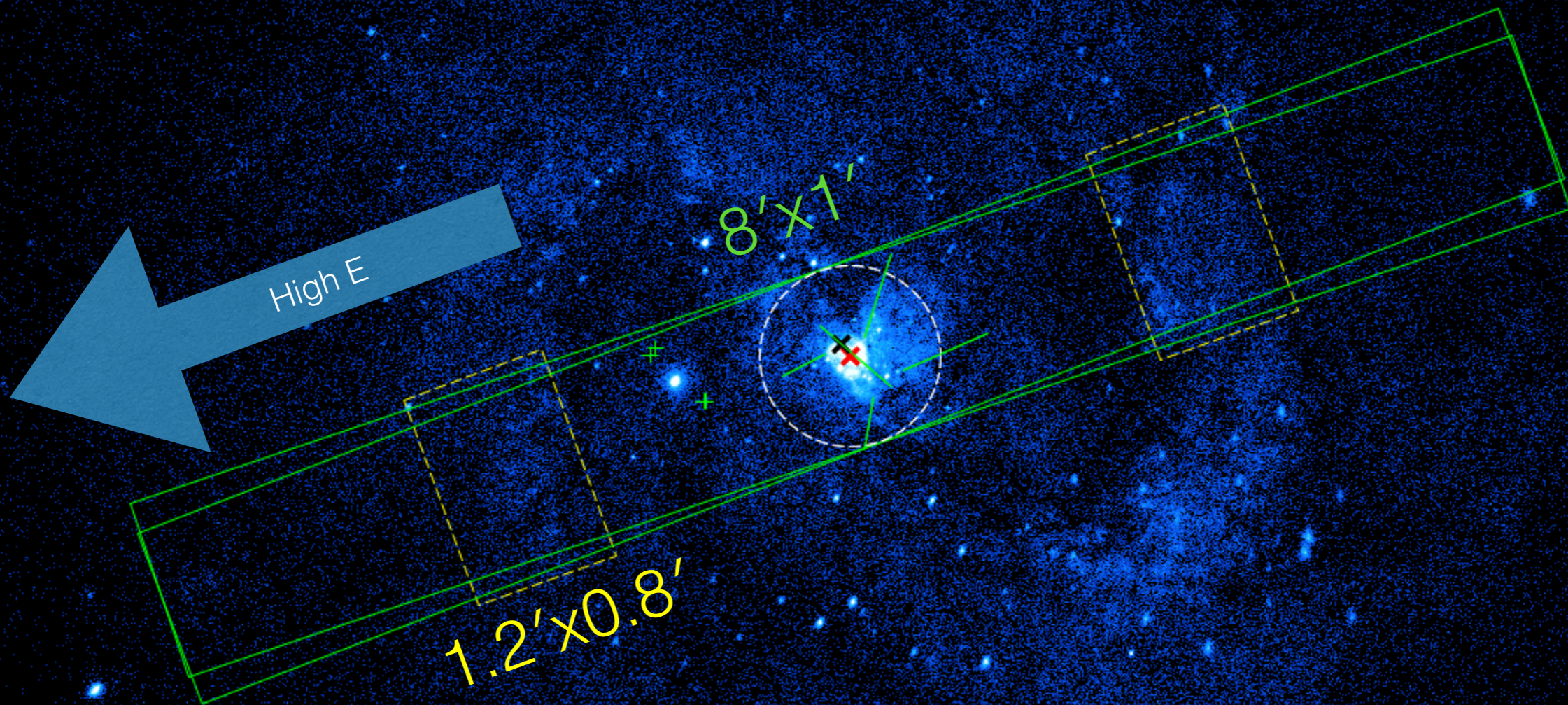


*Bruna+2022; MUSE*

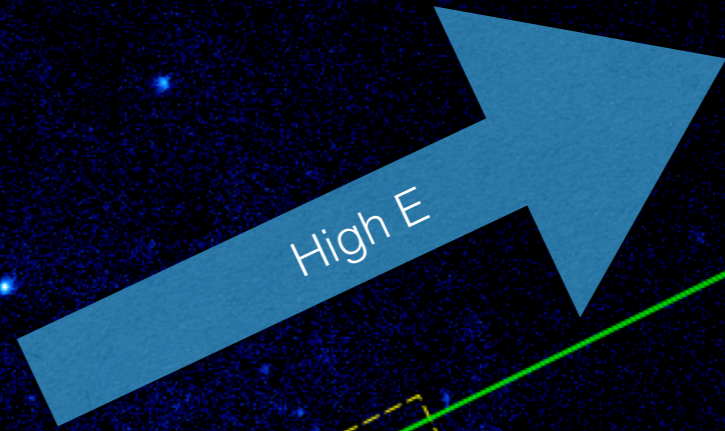


*Chi et al. in preparation; Chandra*

left: PA  $\sim 110^\circ$

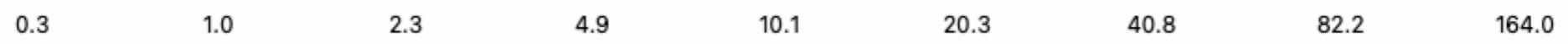
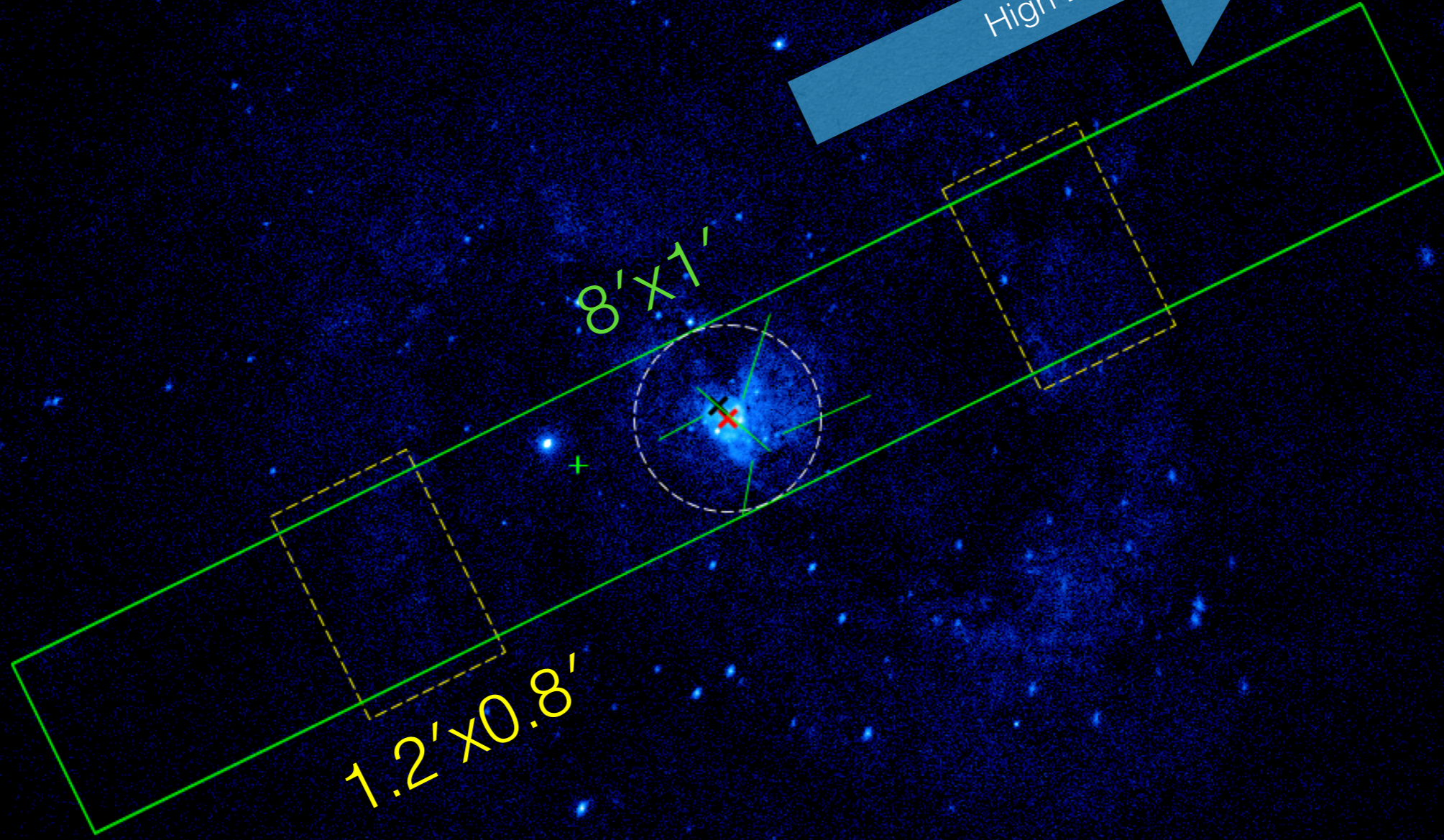


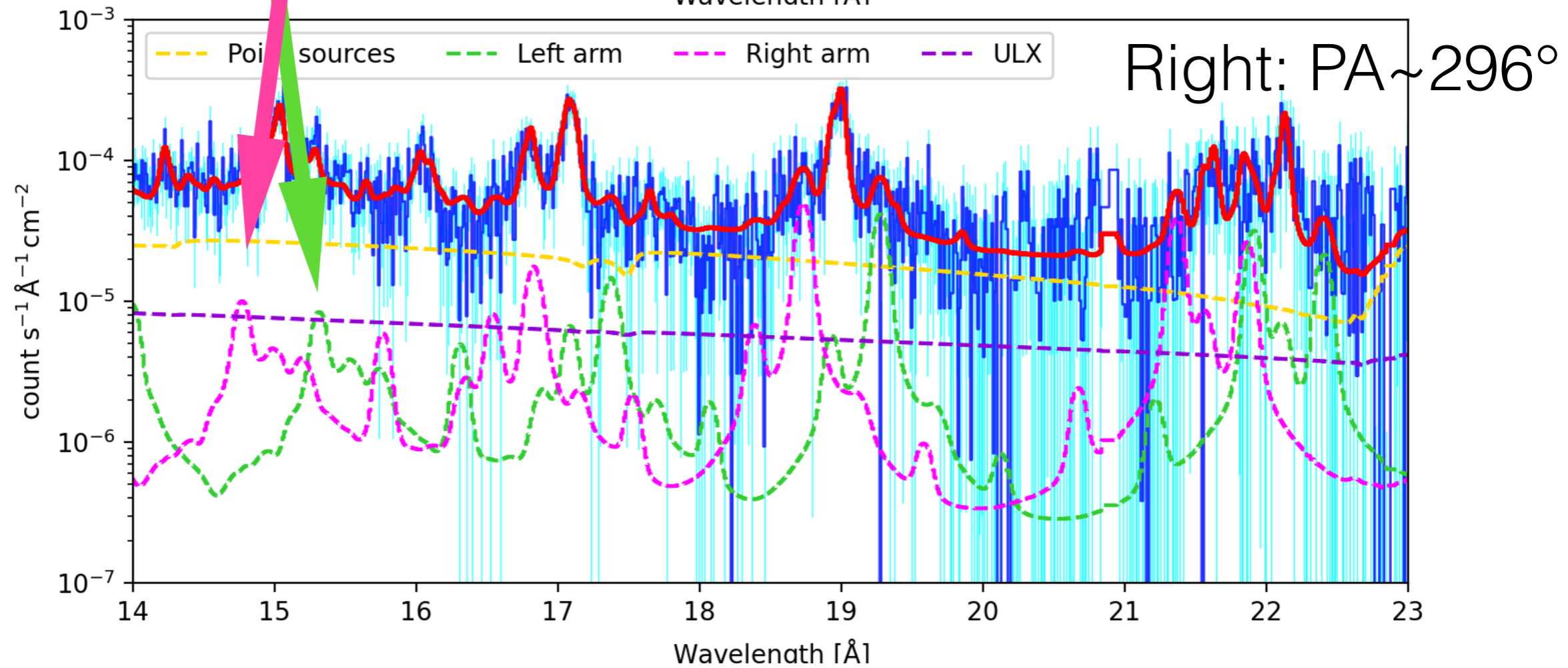
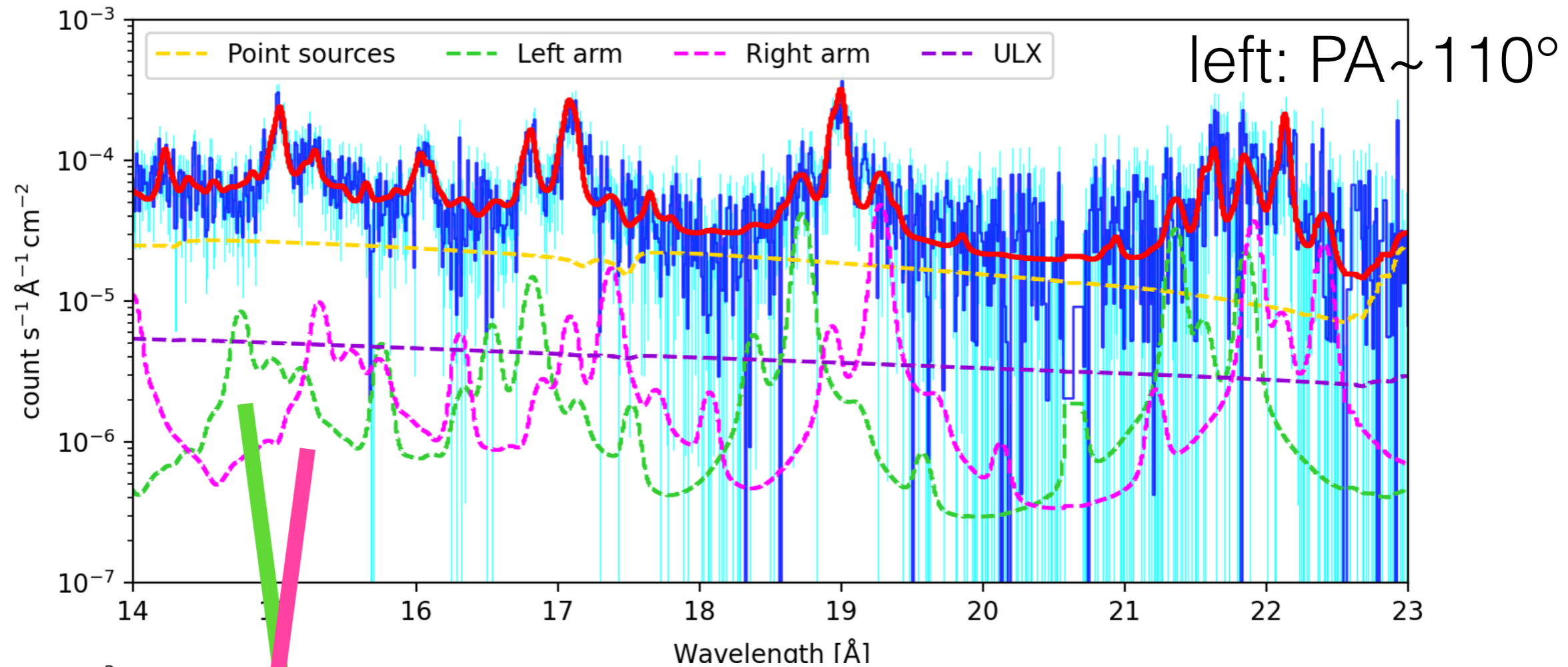
Right: PA  $\sim 296^\circ$



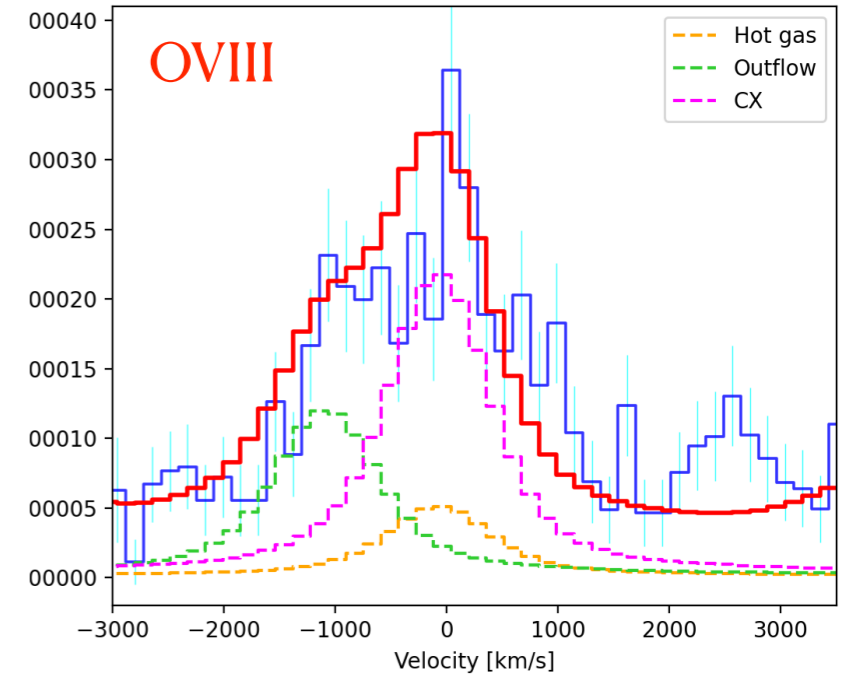
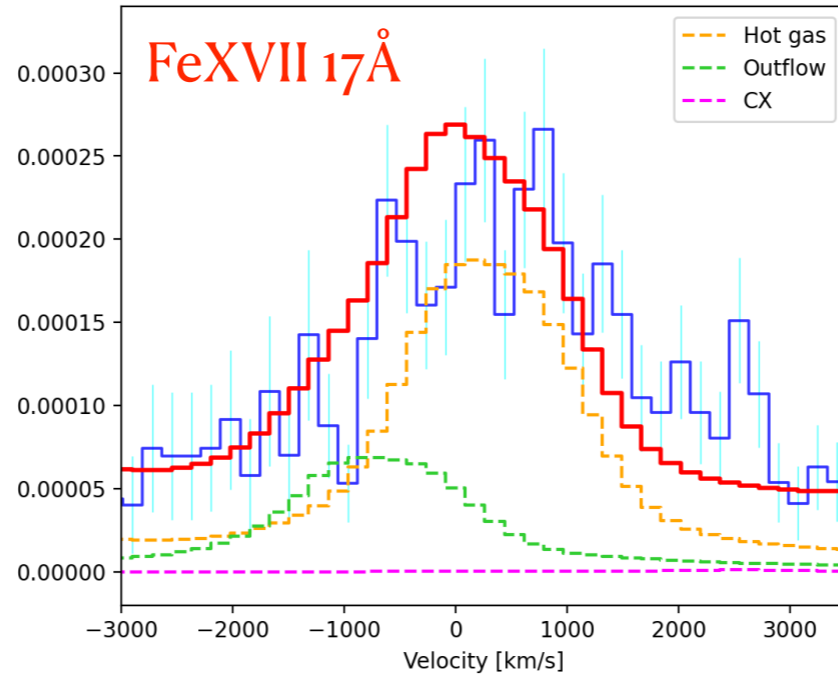
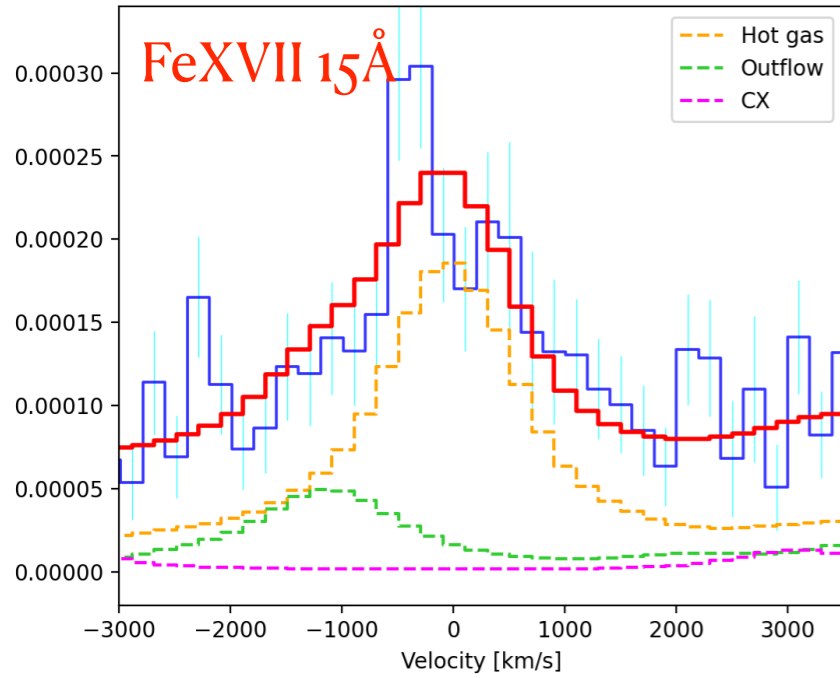
8' x 1'

1.2' x 0.8'

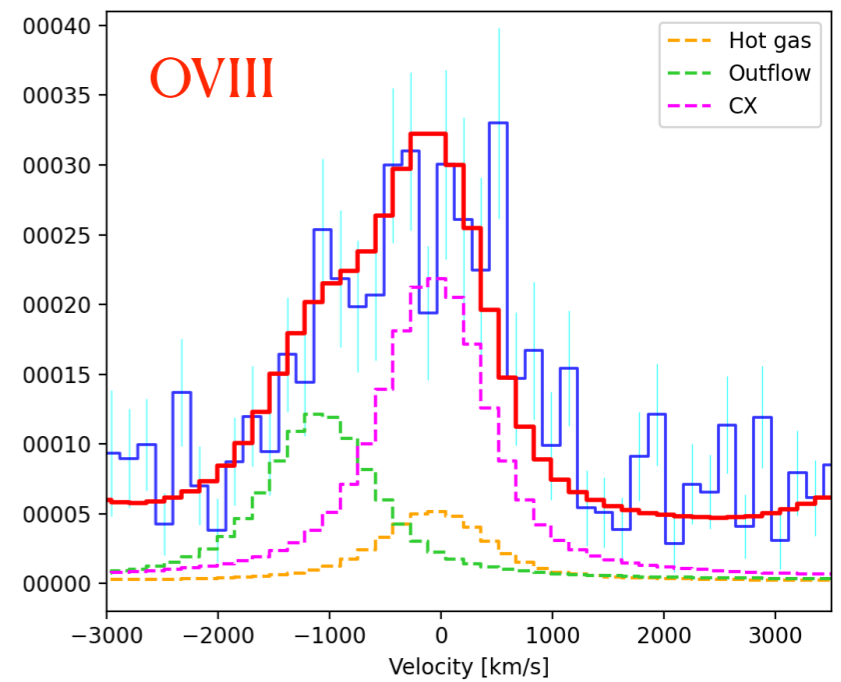
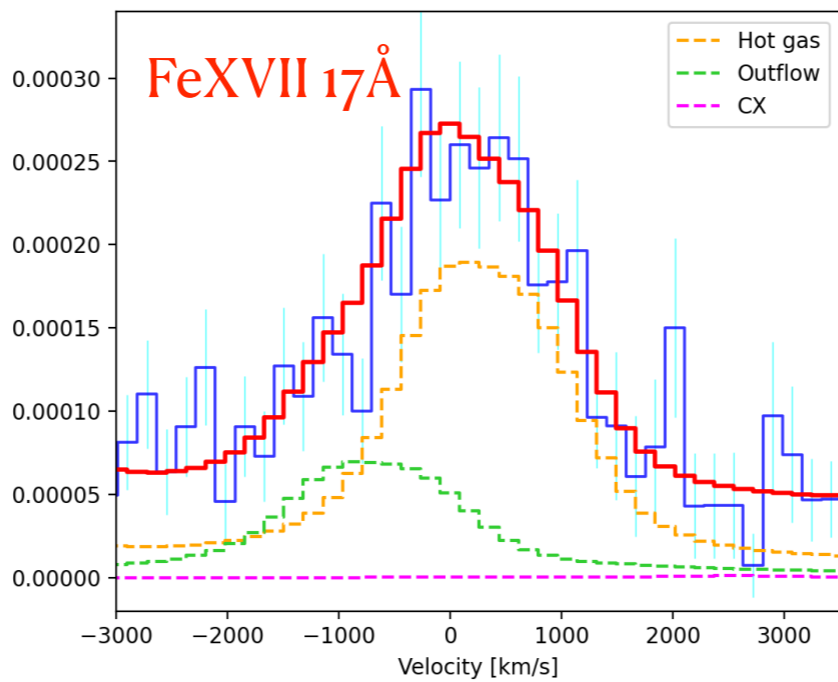
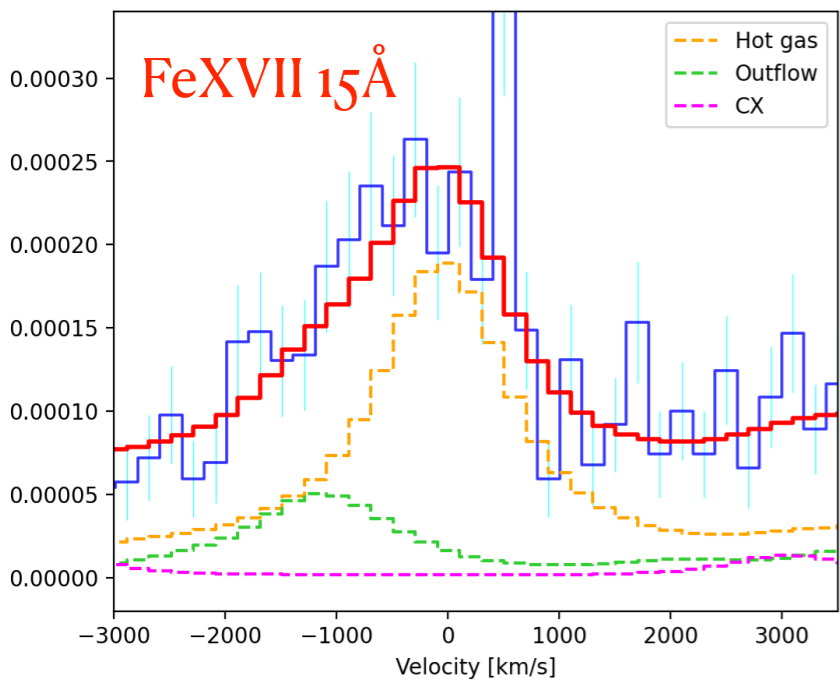




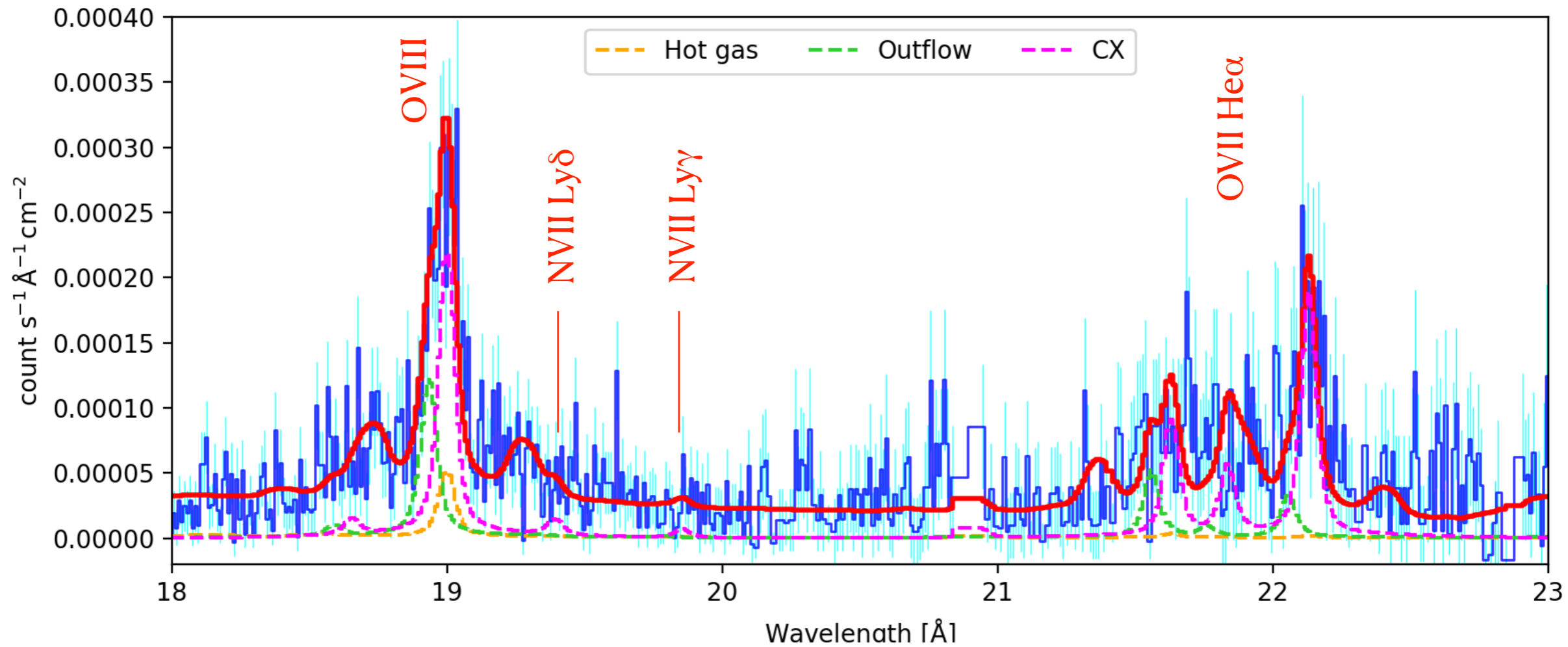
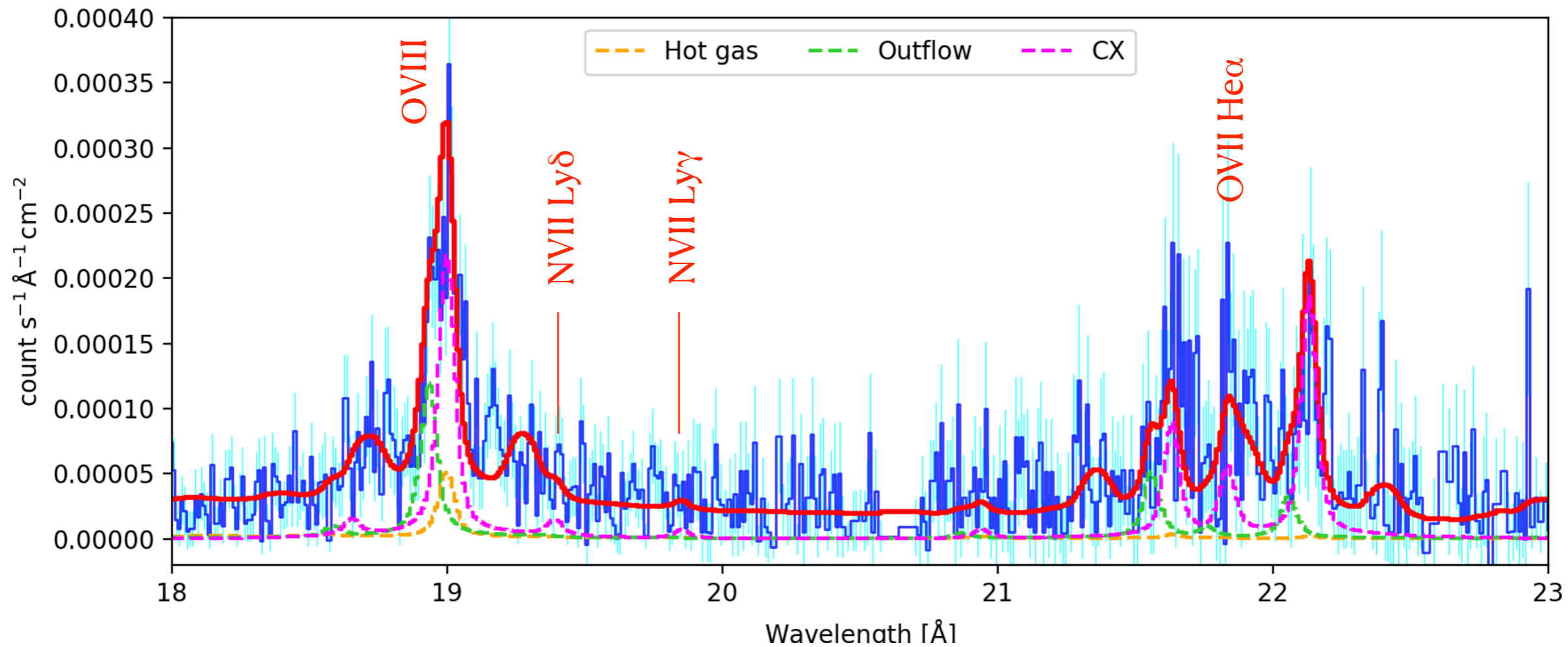
left: PA  $\sim 110^\circ$



Right: PA  $\sim 296^\circ$

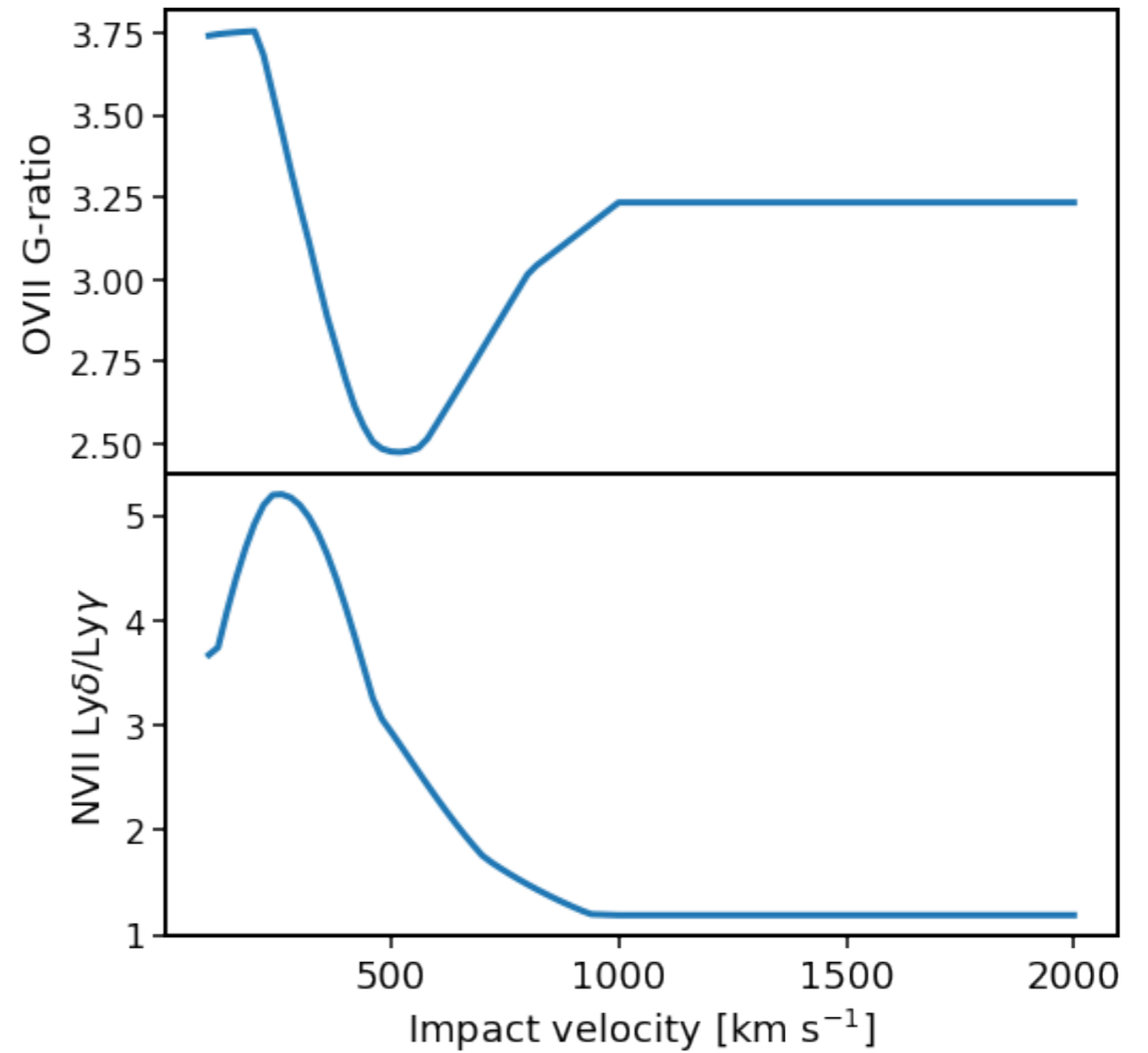


- **Outflowing velocity  $\sim -1000$  km/s.**
- **Soft X-ray emission is from the outflowing hot gas.**

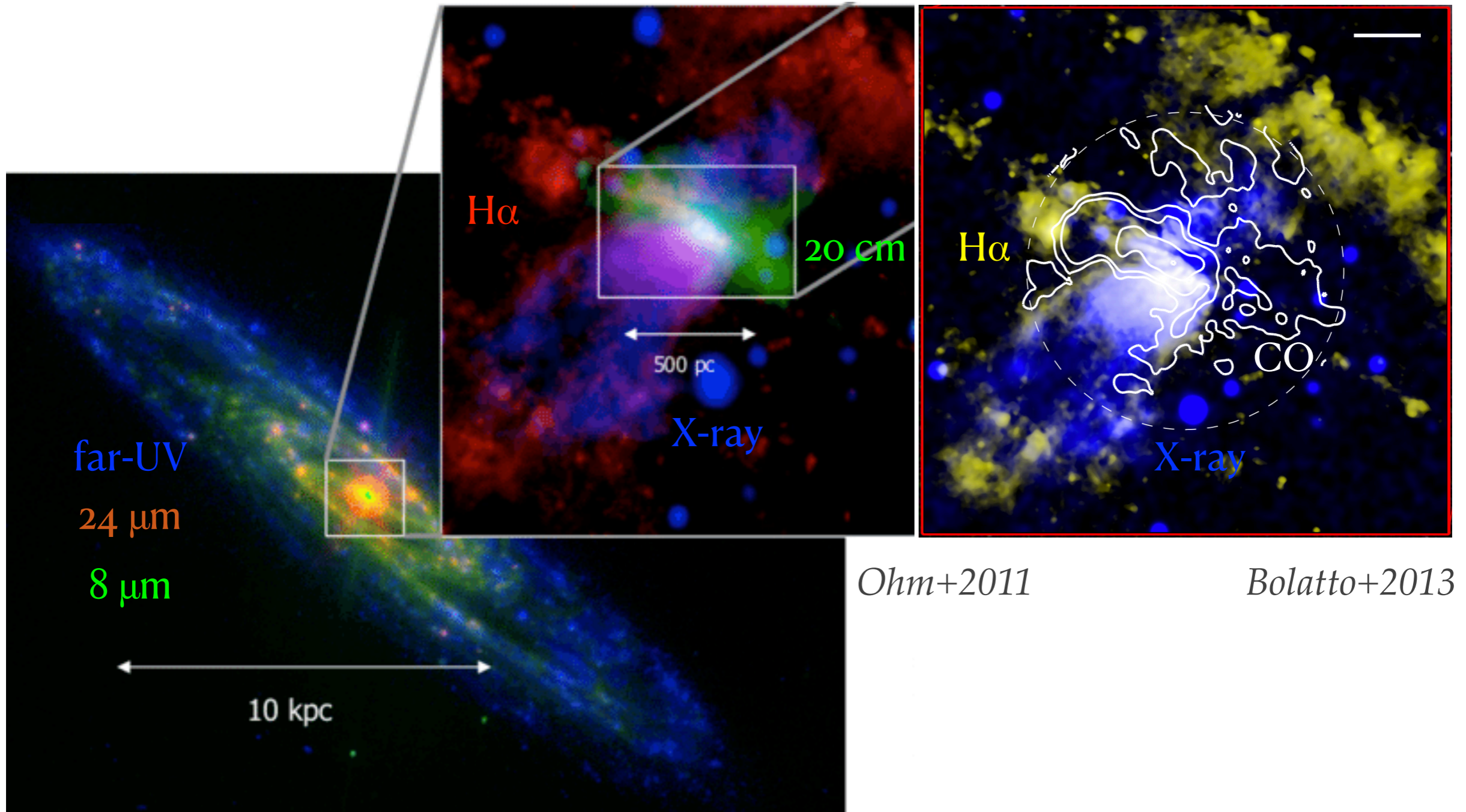


Wait~

Interplay  
Velocity



# NGC253



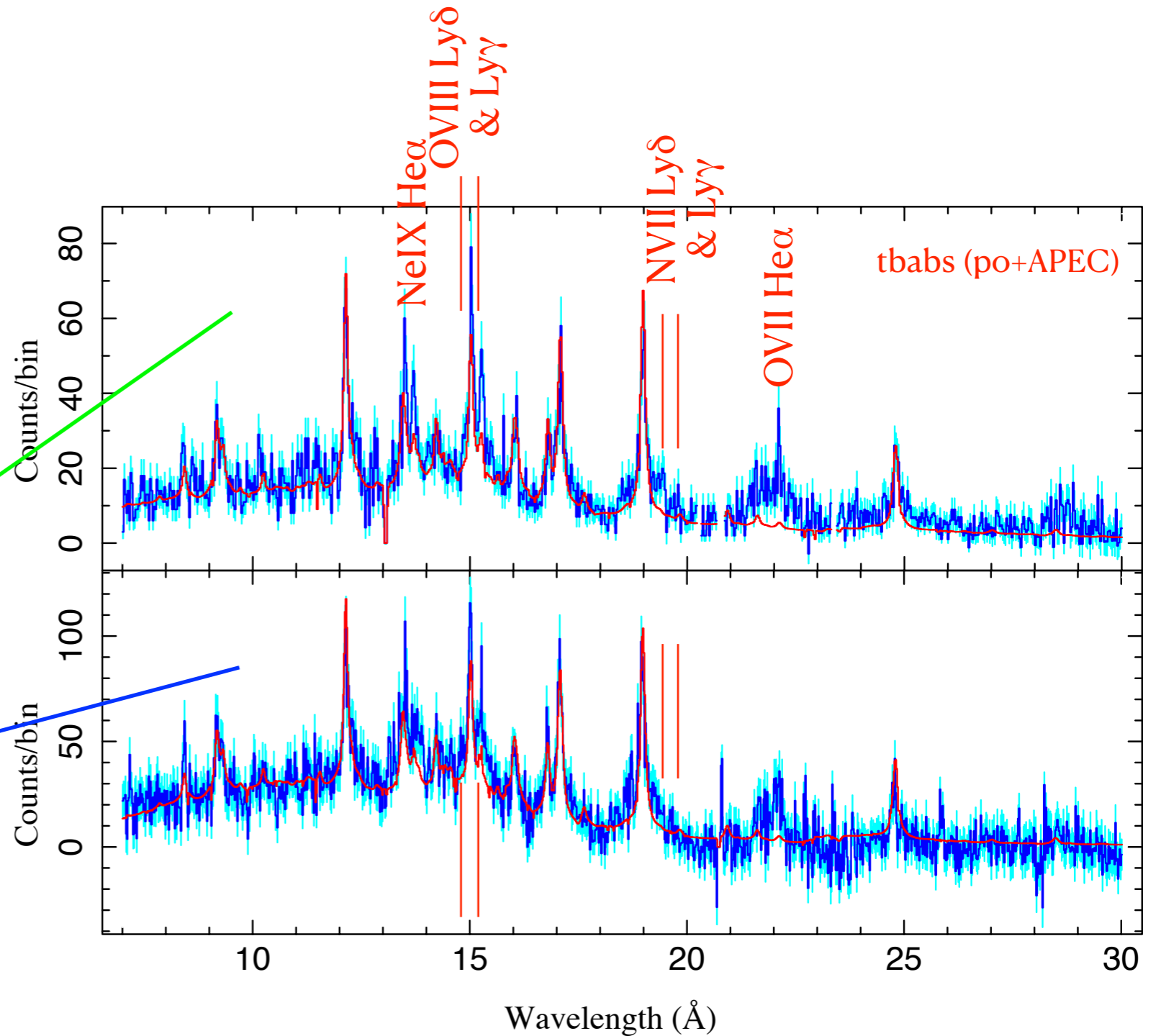
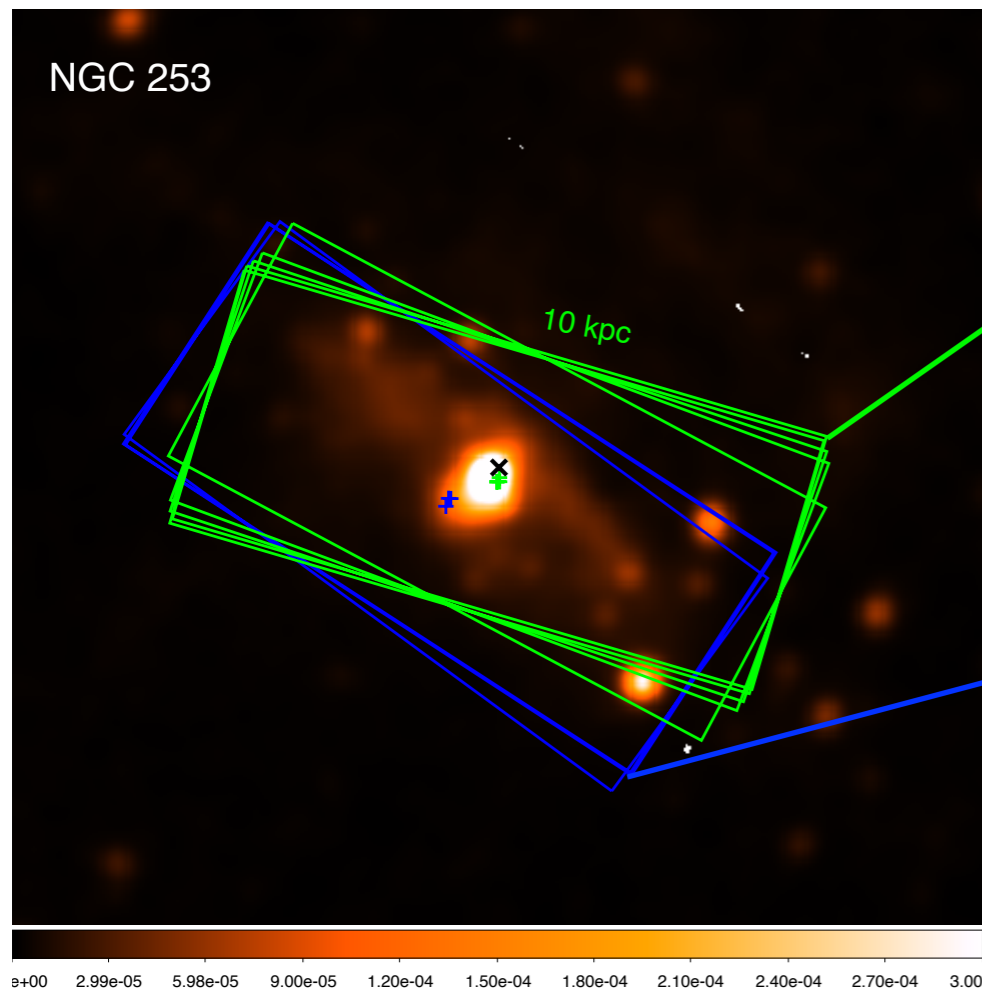
*Ohm+2011*

*Bolatto+2013*

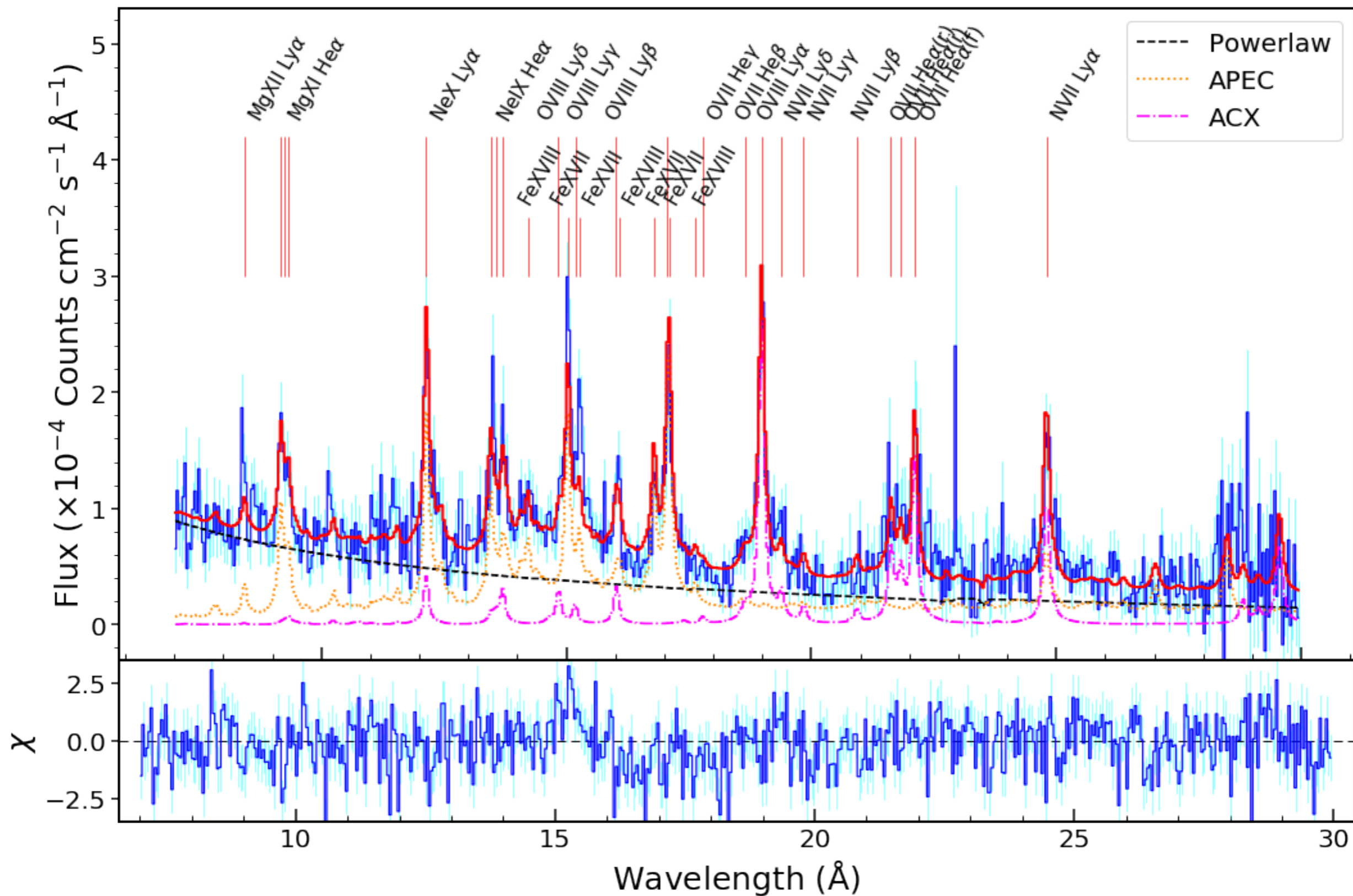


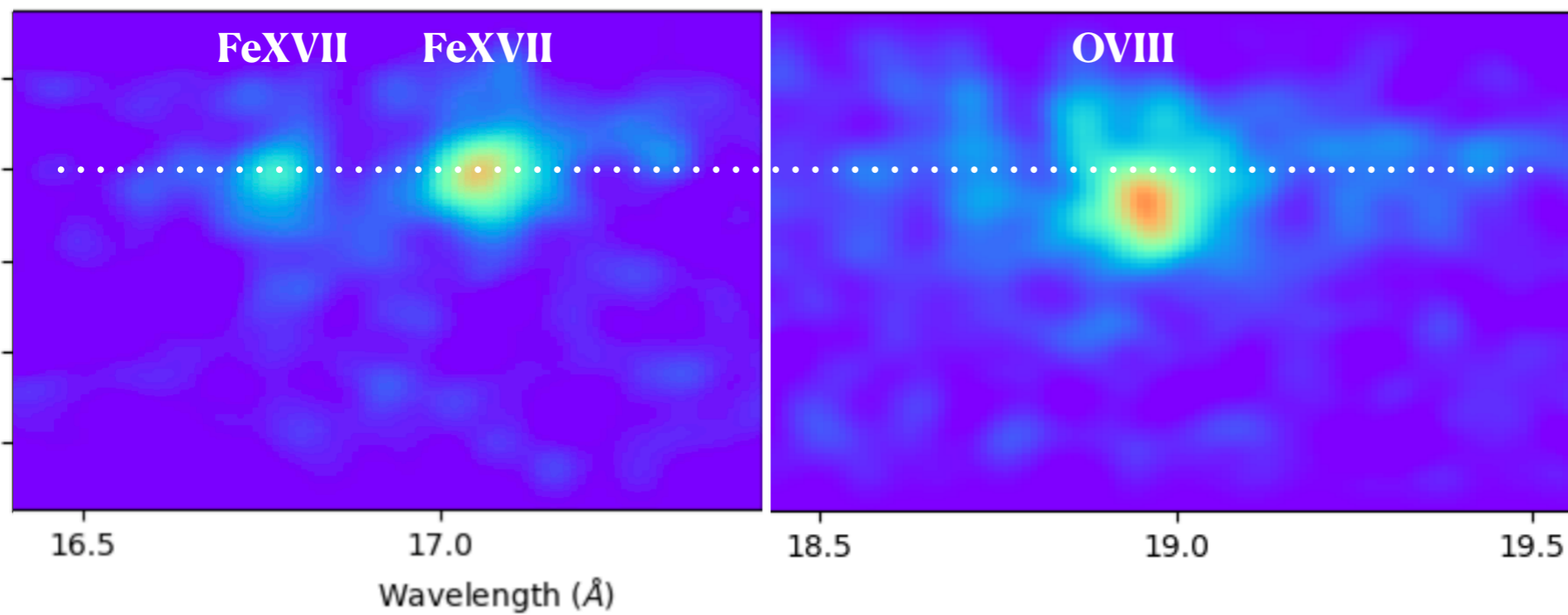
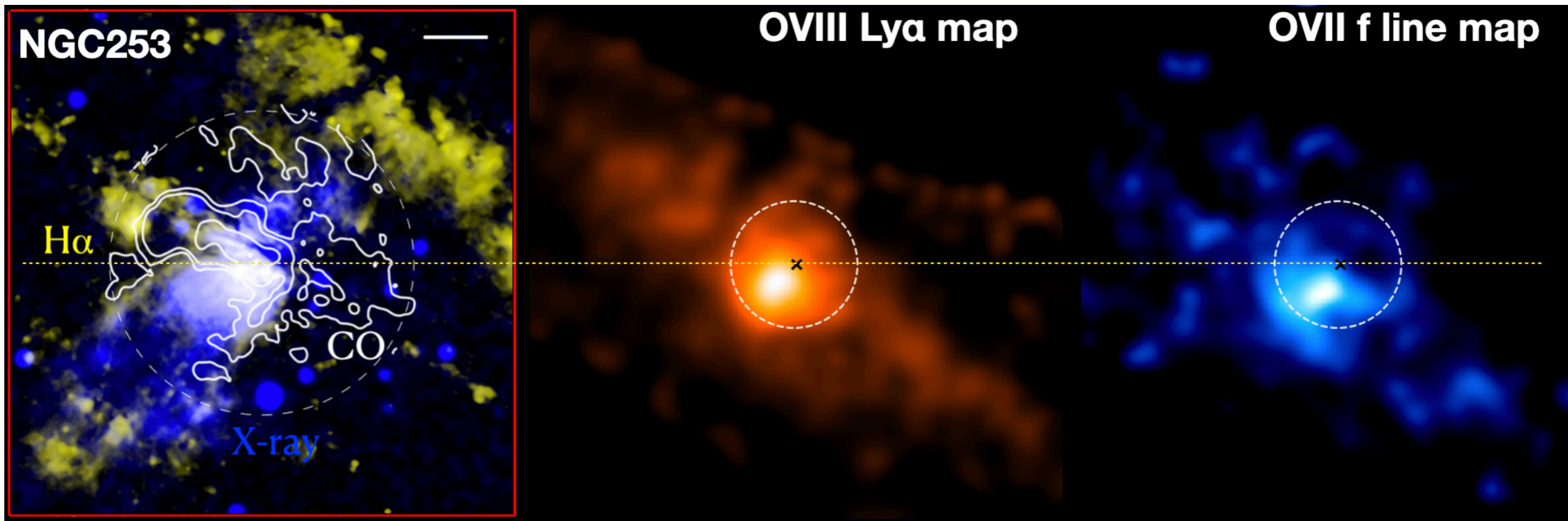
# Opposite dispersion directions

- Blue & Green
- $\text{NVII Ly}\delta > 5\sigma$



# preliminary



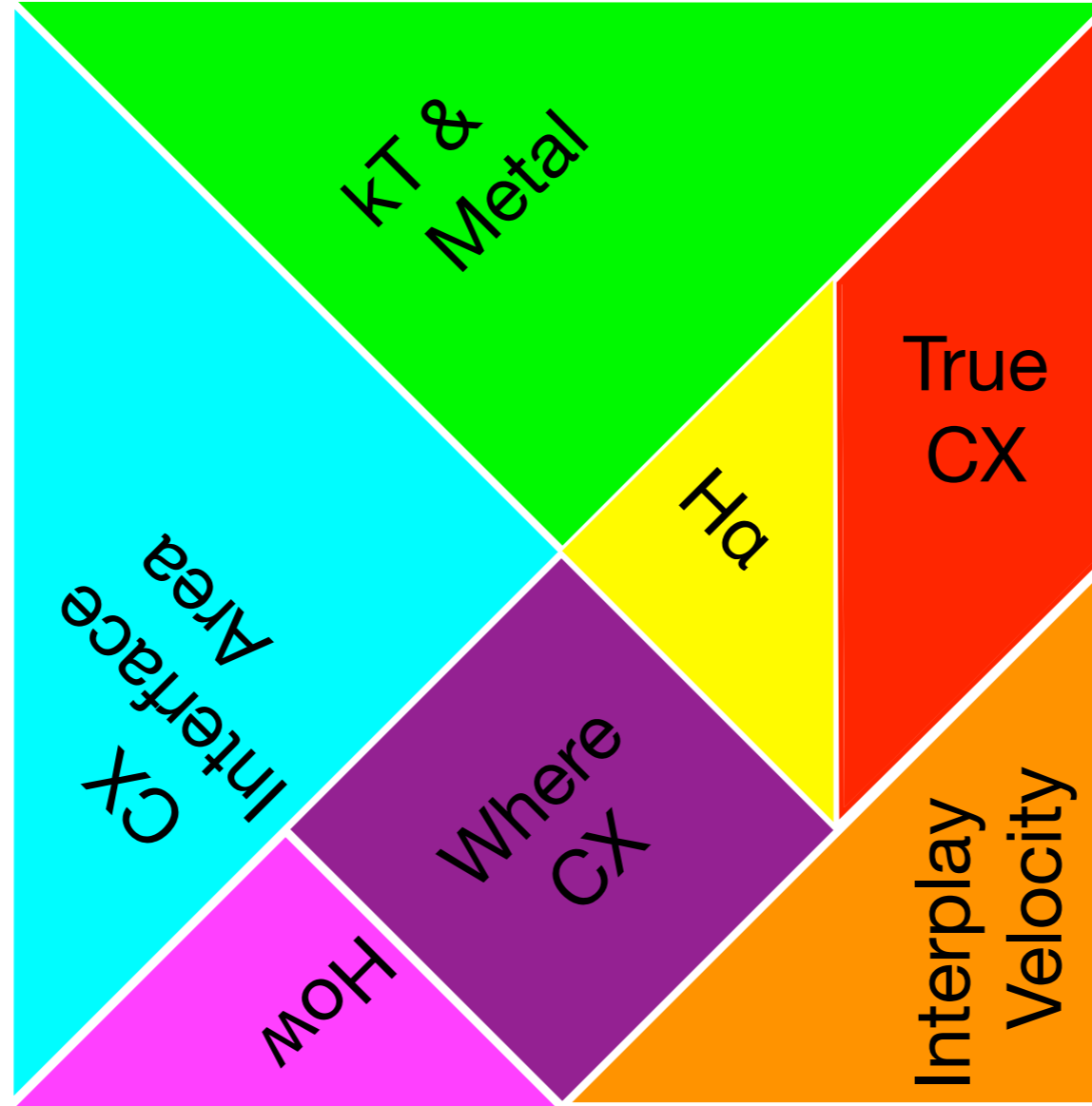


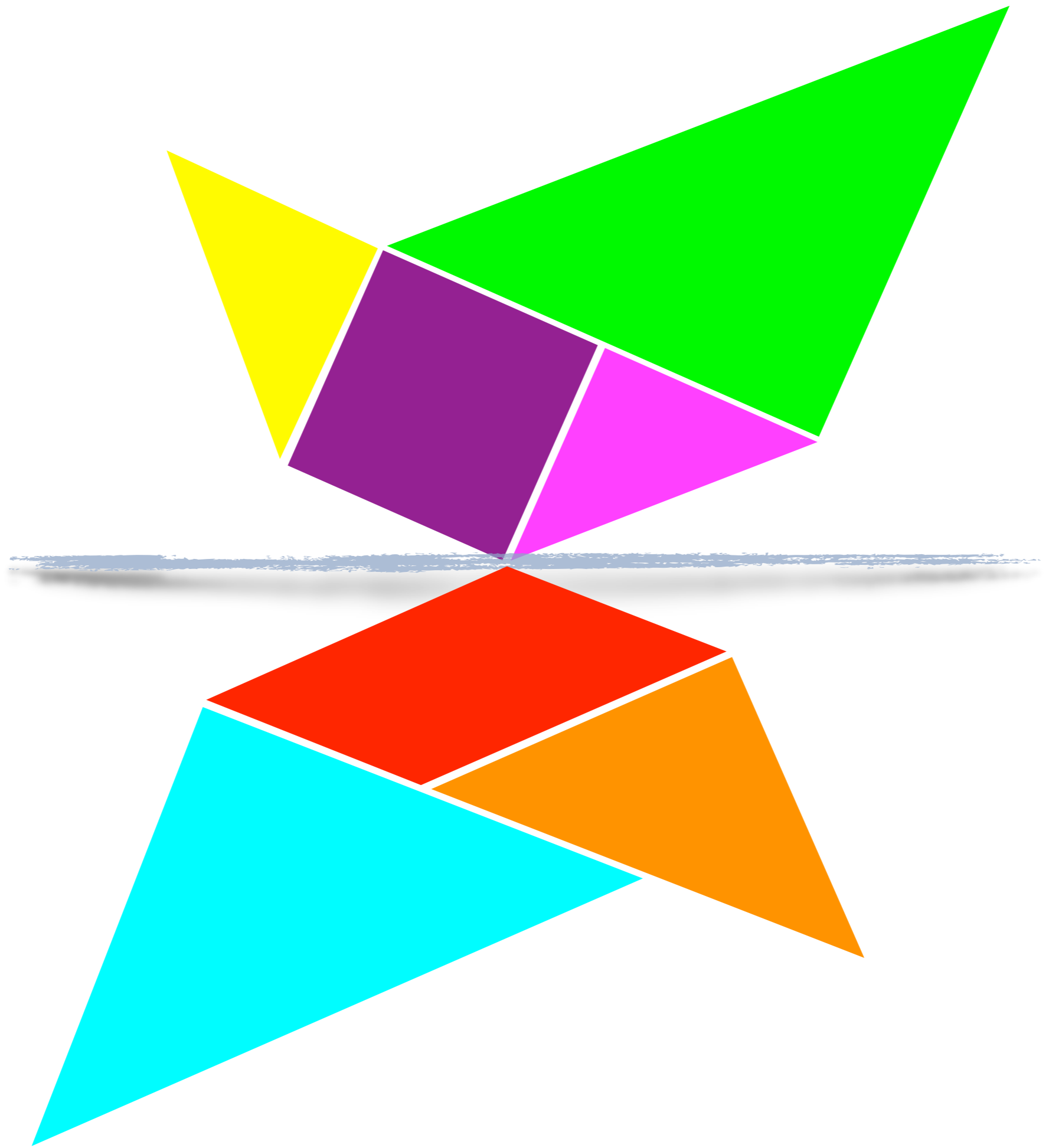
- Outflow blocked
- Low-velocity CX



True  
CX

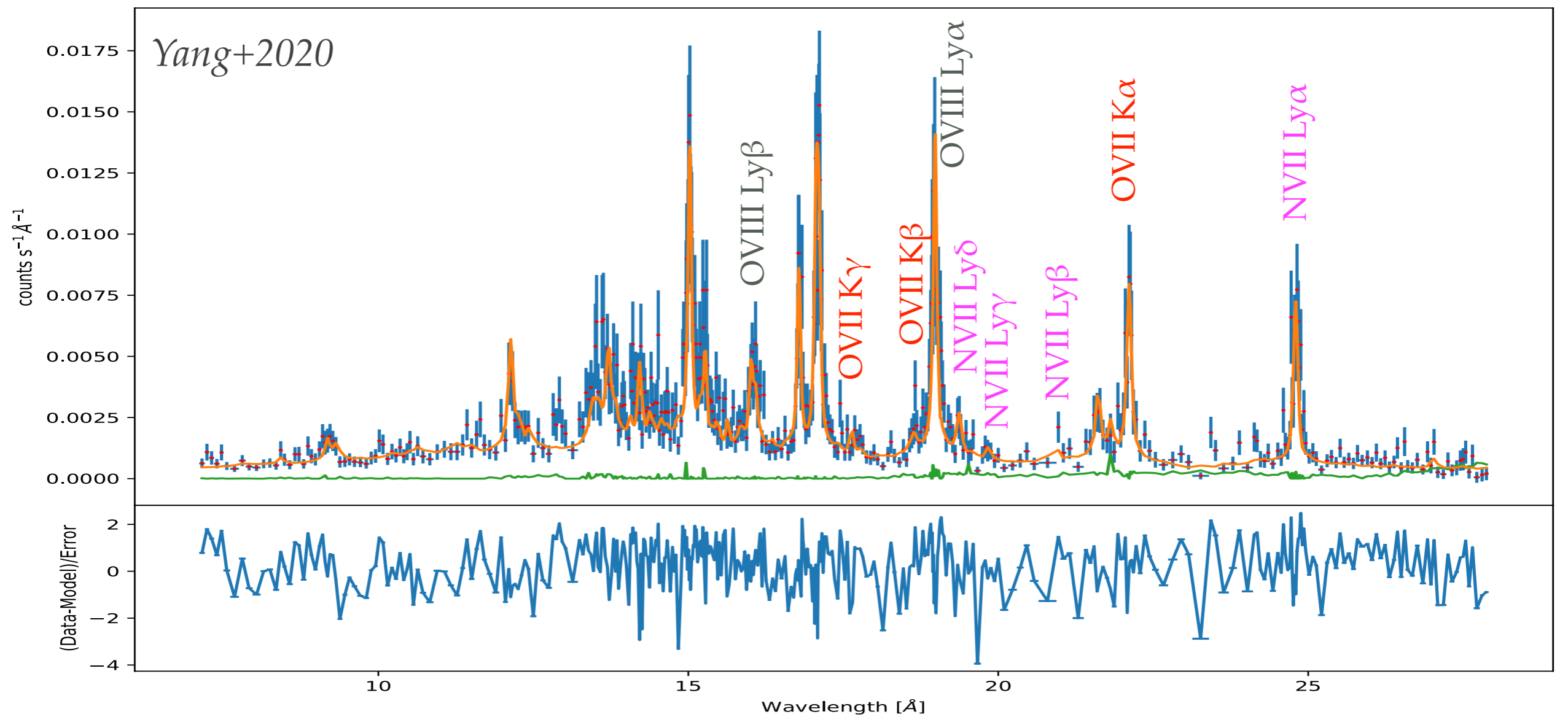
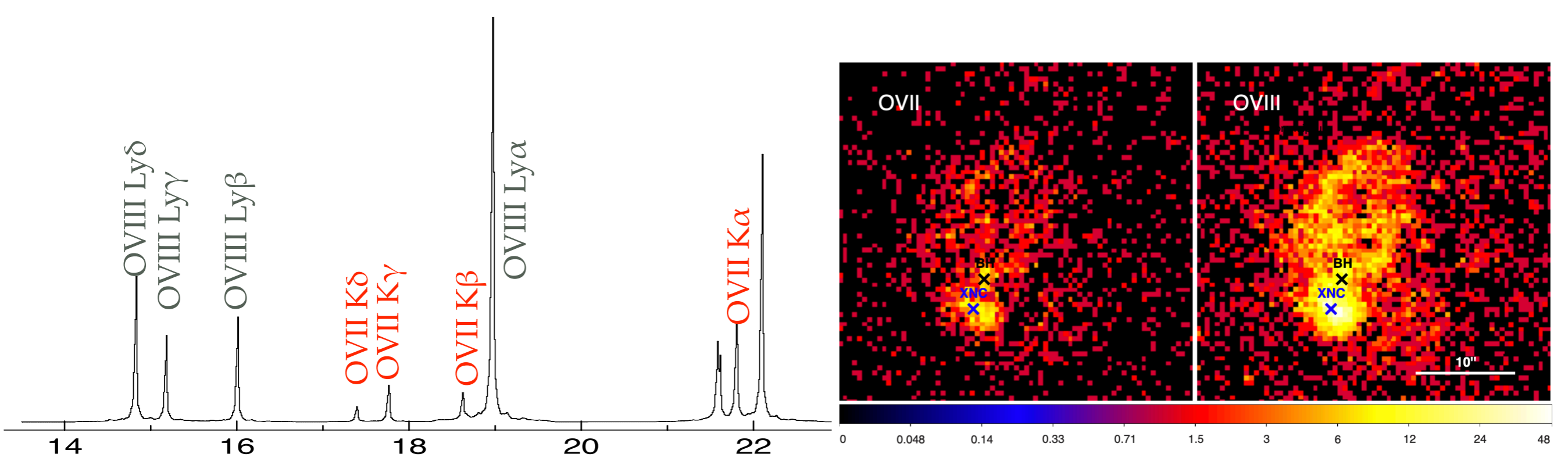
# Tangram







Credit: NASA & ESA

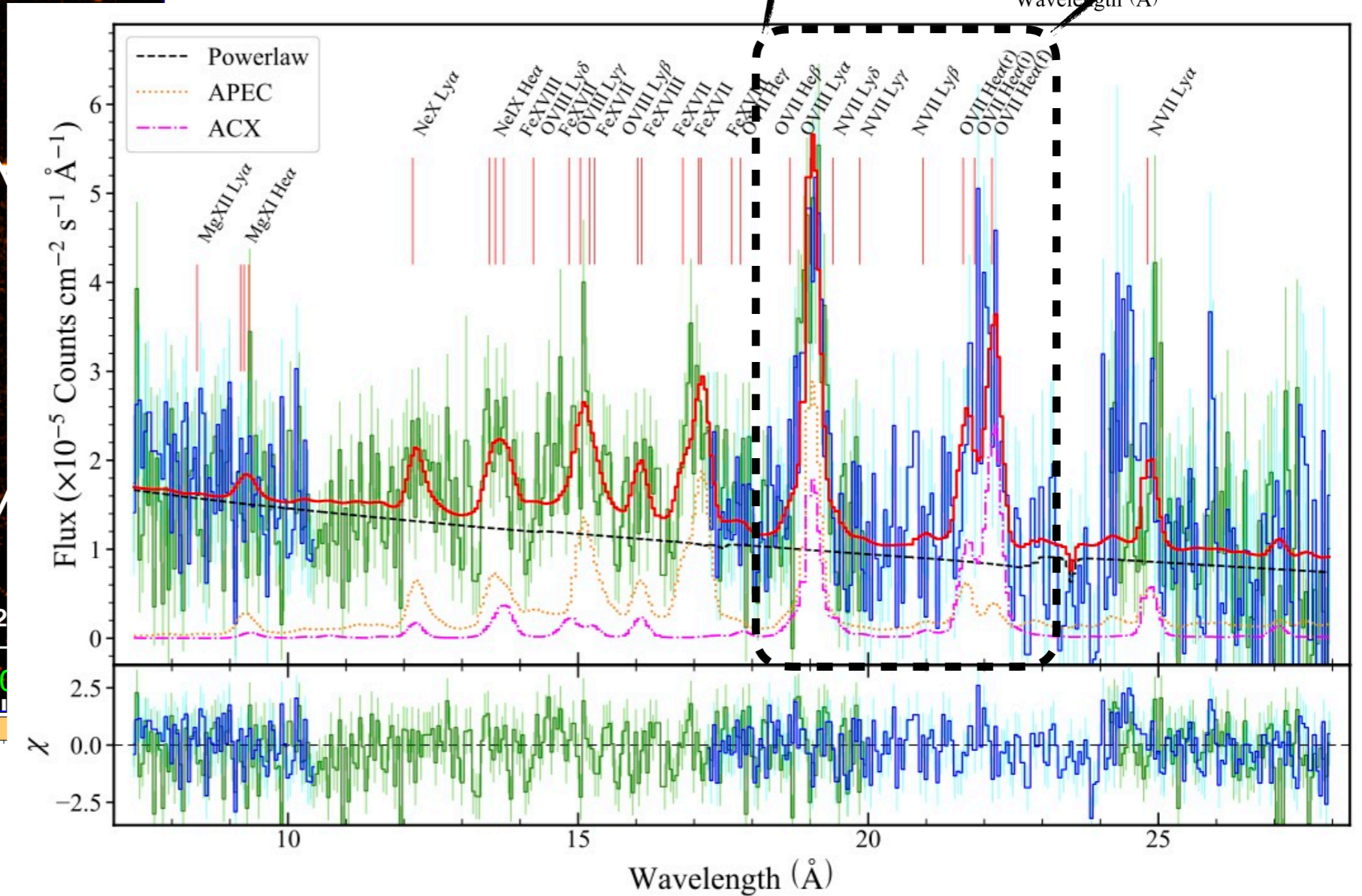
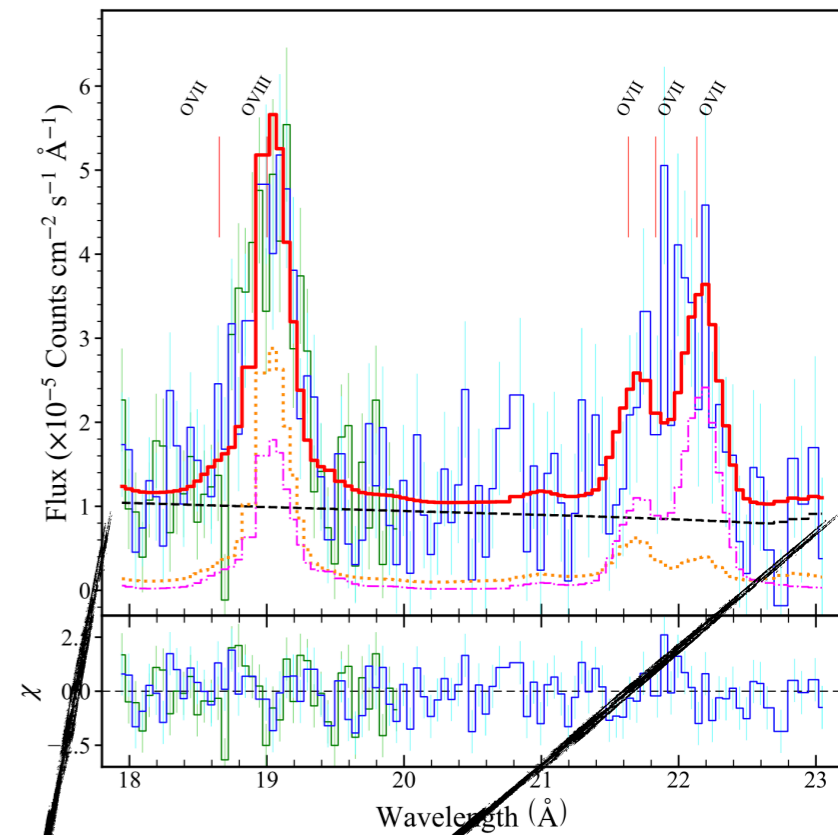
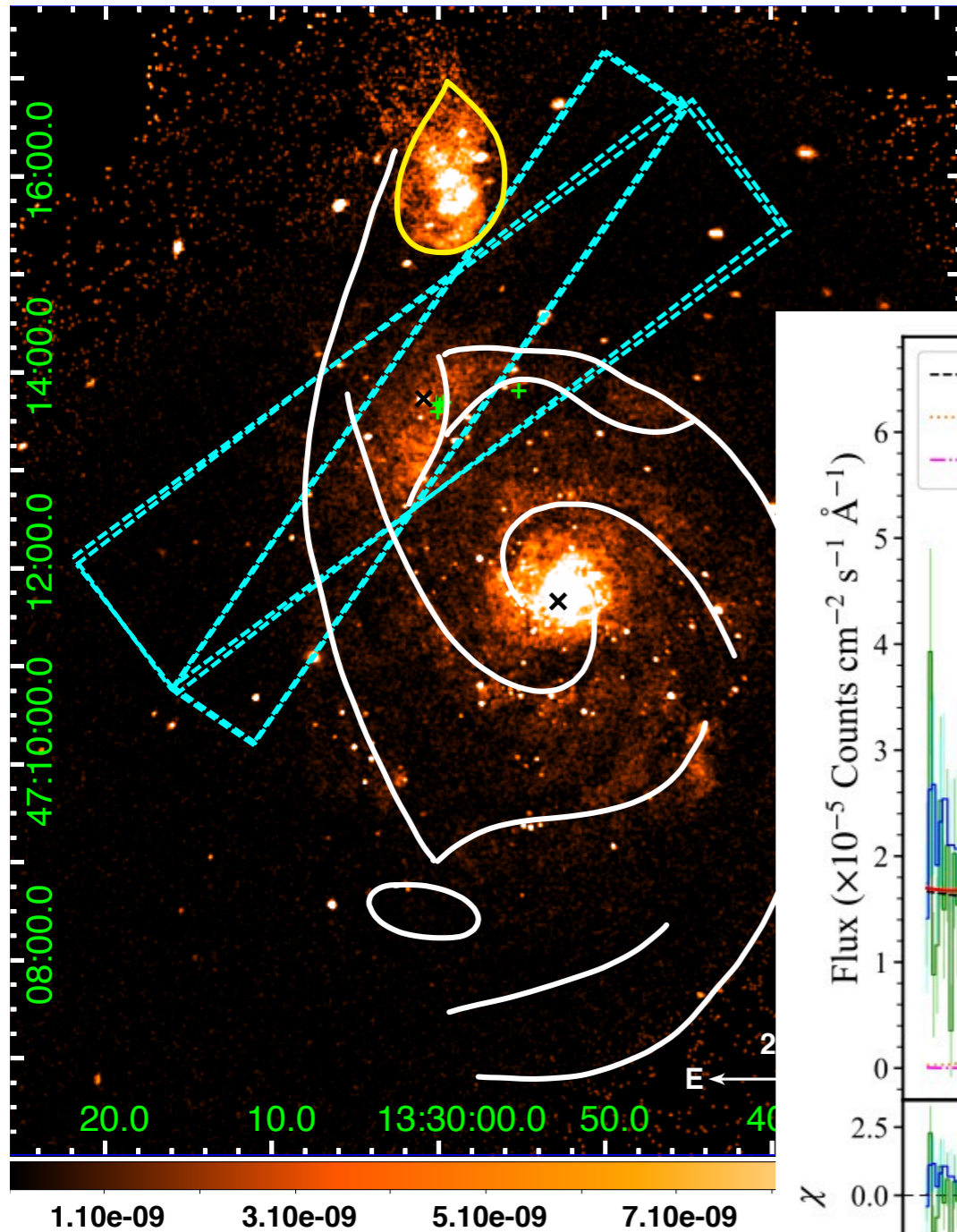




# CX on the disk

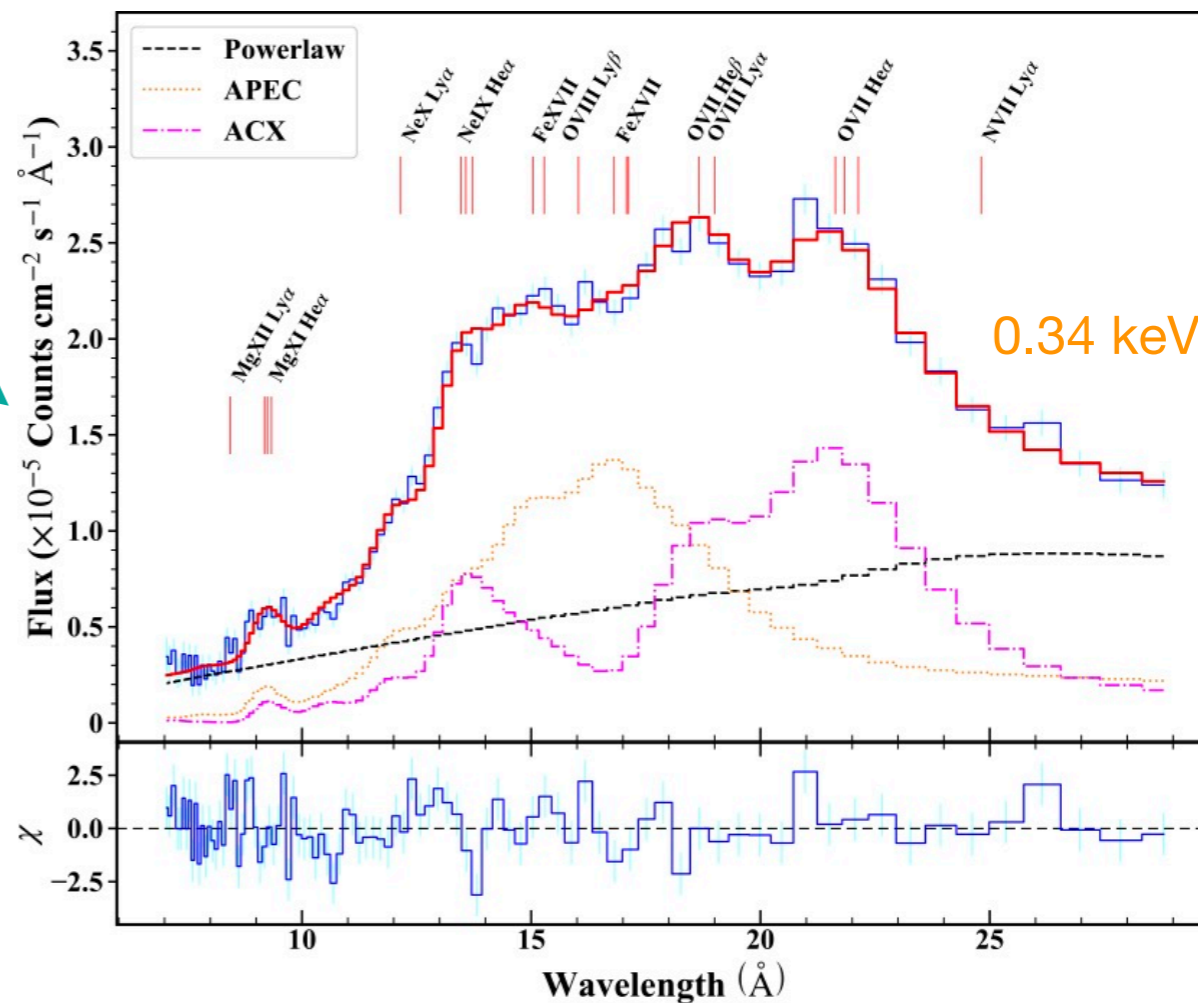
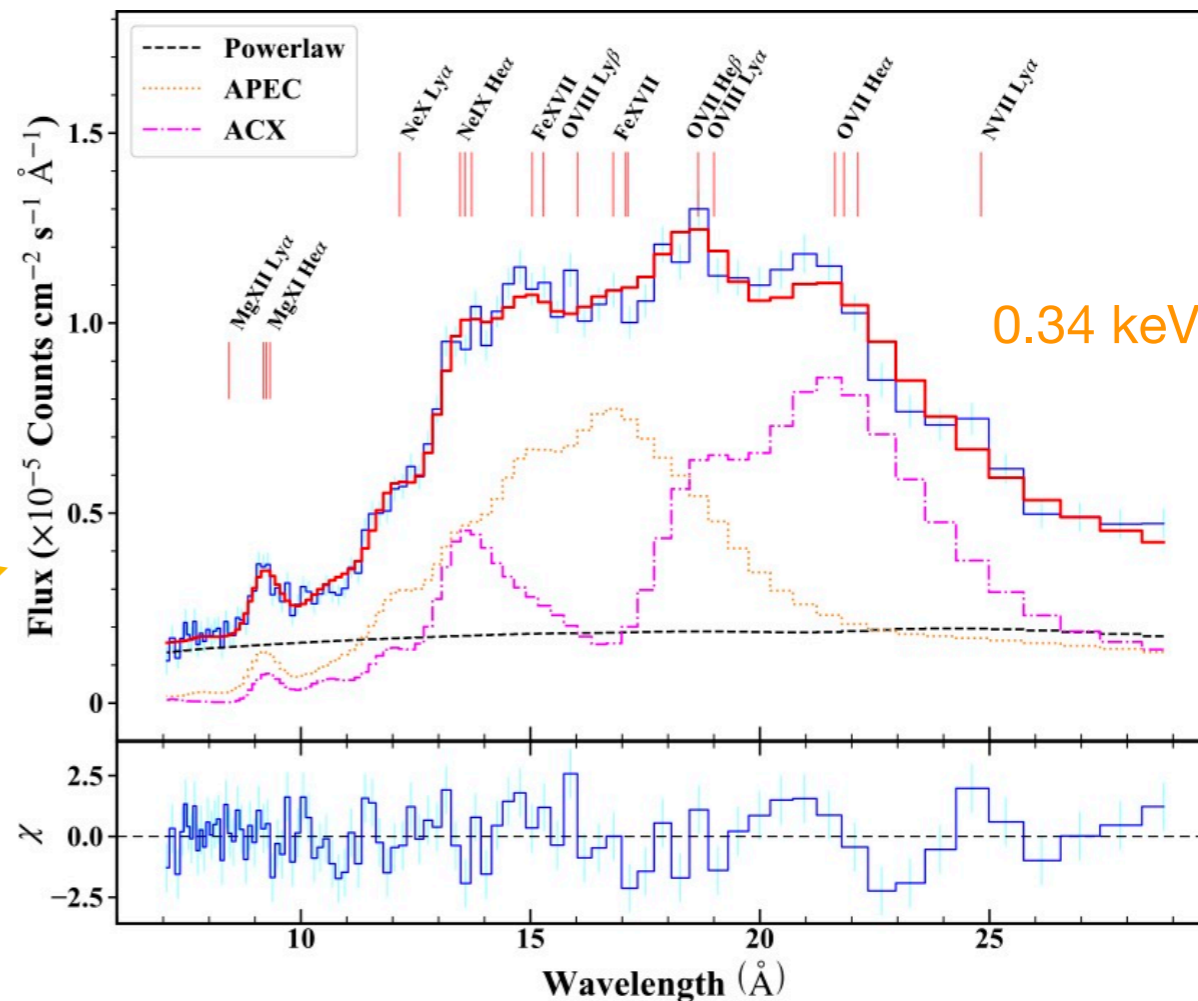
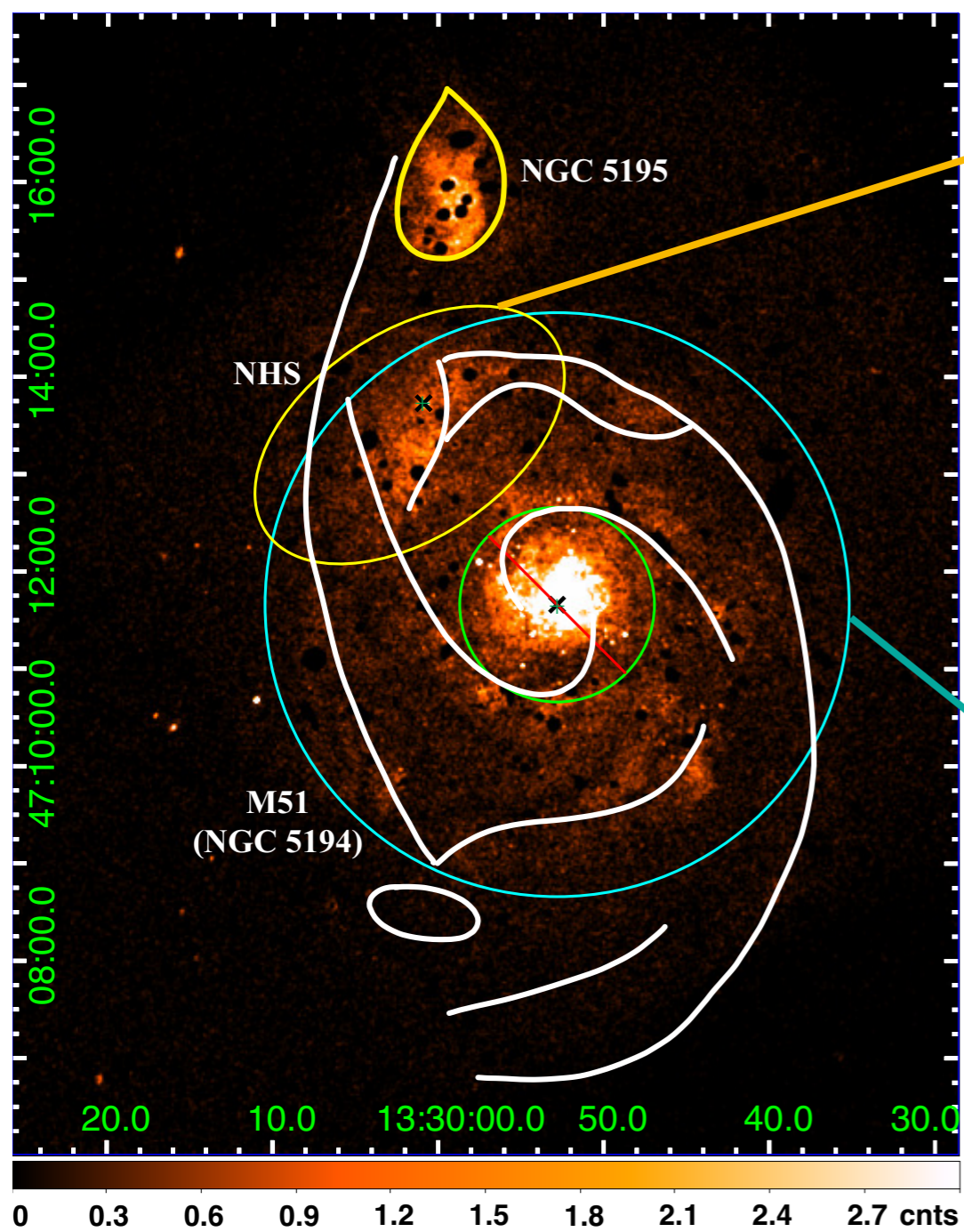
OVII G-ratio:  
 $(f+i)/r$   
 $\sim 3.2^{+6.9}_{-1.5}$

**Not Photoionization,  
 not NEI, not RS**



Zhang+2022

# Chandra/ACIS spectra for the entire disk



**STEPHAN'**

**Chandra X-ray**

**NGC 7319**

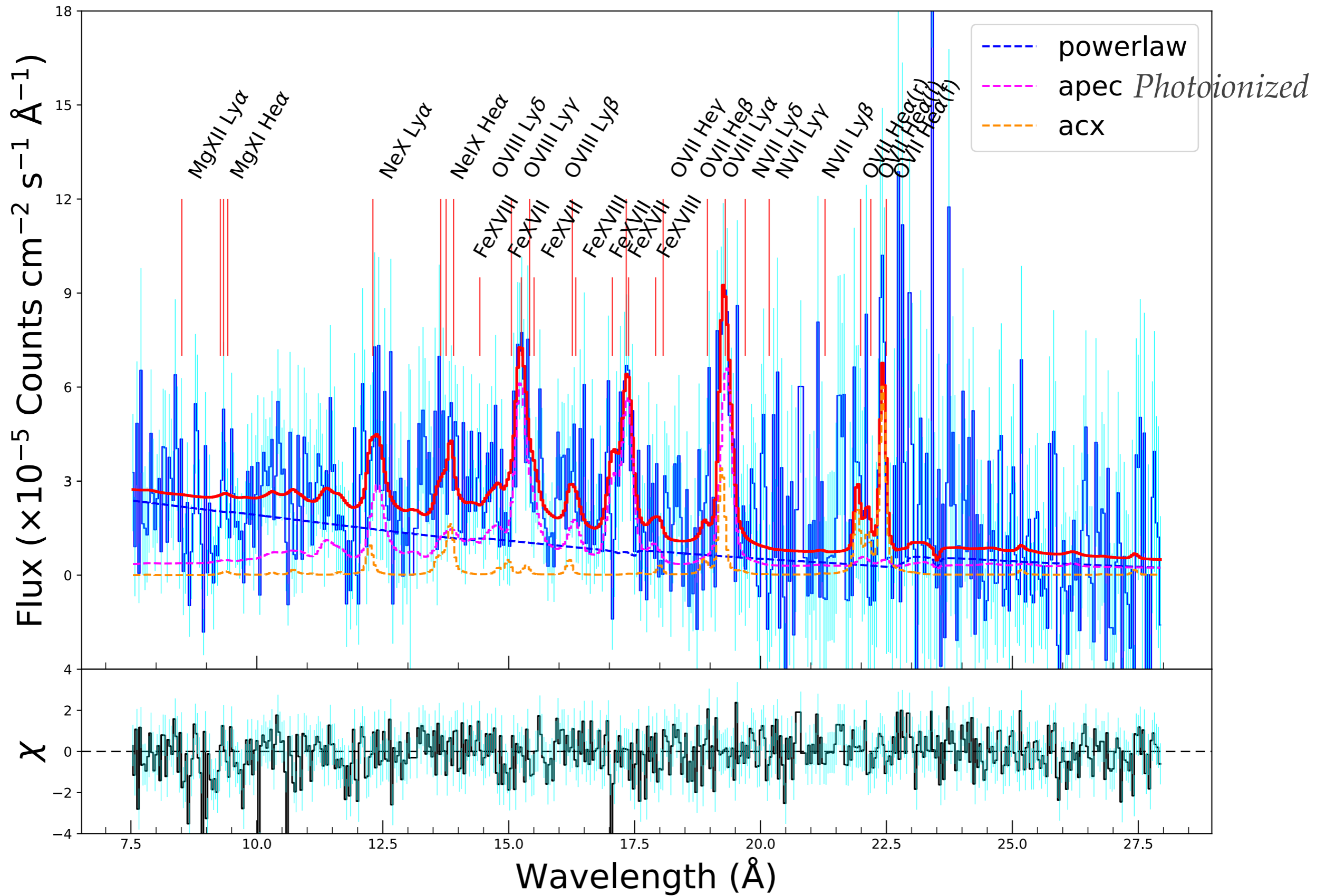
**NGC 7318 b&a**

**NGC 7320**

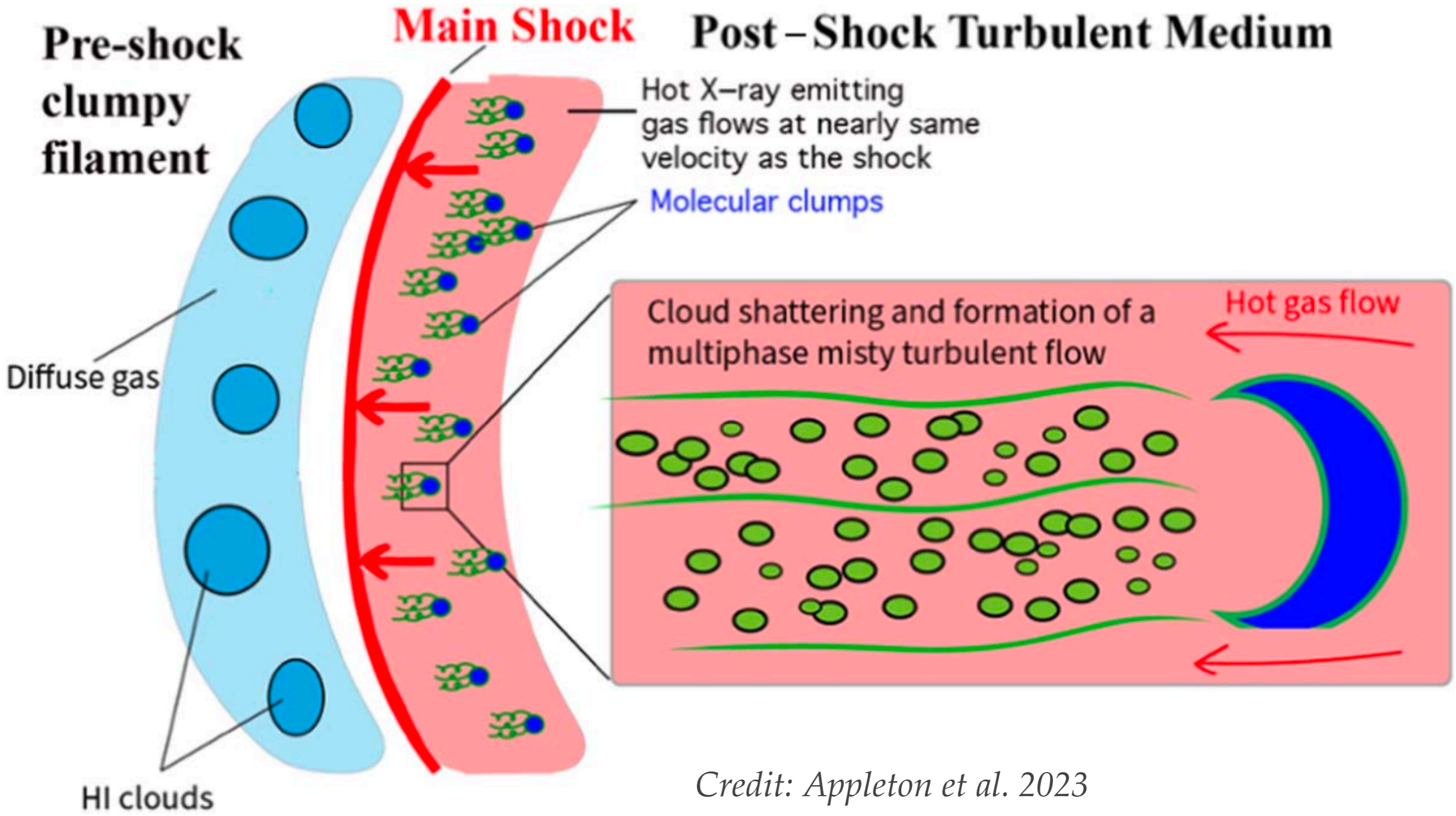
**NGC 7317**

RA.HARVARD.EDU

Credit: JWST/NIRCAM



*Mao et al. in preparation*



*Credit: Appleton et al. 2023*

# Summary

- CX may everywhere in the galaxy: center, disk, superwind, outside.
- CX helps constrain properties of both the hot and the cold gas.
- Sometimes, the hot gas interacts with the cold ISM severely, which consumes the hot gas quickly.